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# Learning from model grids: Tracers of the ionization fraction in the ISM

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**Abstract.** The ionization fraction in neutral interstellar clouds is a key physical parameter controlling multiple physical and chemical processes, and varying by orders of magnitude from the UV irradiated surface of the cloud to its cosmic-ray dominated central regions. Traditional observational tracers of the ionization fraction, which mostly rely on deuteration ratios of molecules like HCO<sup>+</sup>, suffer from the fact that the deuterated molecules are only detected in a tiny fraction of a given Giant Molecular Cloud (GMC). In [1], we propose a *machine learning*-based, semi-automatic method to search in a large dataset of astrochemical model results for new tracers of the ionization fraction, and propose several new tracers relevant in different ranges of physical conditions.

## 1 Context

The ionization fraction  $x(e^-)$  is a key physical parameter controlling several aspects of the physics and chemistry of neutral interstellar clouds. Fast ion-neutral reactions are one of the main drivers of interstellar chemistry. The ionization fraction also controls the coupling of the gas to the magnetic field, as well as the charge state of grains and thus the efficiency of the photoelectric heating mechanism on grains. Finally, it can affect our interpretation of the observed excitation of high dipole moment molecules such as HCN, which are very sensitive to collisional excitation by electrons. In a typical Giant Molecular Cloud,  $x(e^-)$  can vary in the neutral gas from  $\sim 10^{-4} \text{ cm}^{-3}$  at the UV illuminated surface of the cloud, where carbon is fully ionized, to  $10^{-6} - 10^{-9} \text{ cm}^{-3}$  in the central shielded regions, where cosmic rays maintain a residual source of ionization.

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Usual tracers of the ionization fraction use the molecular deuteration of species such as  $\text{HCO}^+$  [3] or  $\text{N}_2\text{H}^+$  [4] that indirectly trace the  $\text{H}_2\text{D}^+/\text{H}_3^+$  ratio, whose enrichment at low temperature is limited by electronic dissociative recombination and which is thus sensitive to the ionization fraction. However deuteration-based tracers of the ionization fraction suffer from several limitations. They are sensitive to the ortho-para ratio of  $\text{H}_2$ , itself a parameter difficult to estimate observationally. They are also relatively insensitive at very low ionization fractions, yielding estimates with large uncertainties. More importantly, these deuterated species are typically only detected in dense cores, which represent only small fraction of the observable area of a GMC. In the context of the Orion-B IRAM-30m Large Program [7], which has mapped the Orion B GMC over 5 square degrees with an angular resolution of  $26''$  in the full 71 to 116 GHz band, we thus aim at finding new tracers of the ionization fraction to allow a full scale, unbiased view of the ionization fraction in a GMC, covering all the types of environments constituting a GMC.

## 2 Method

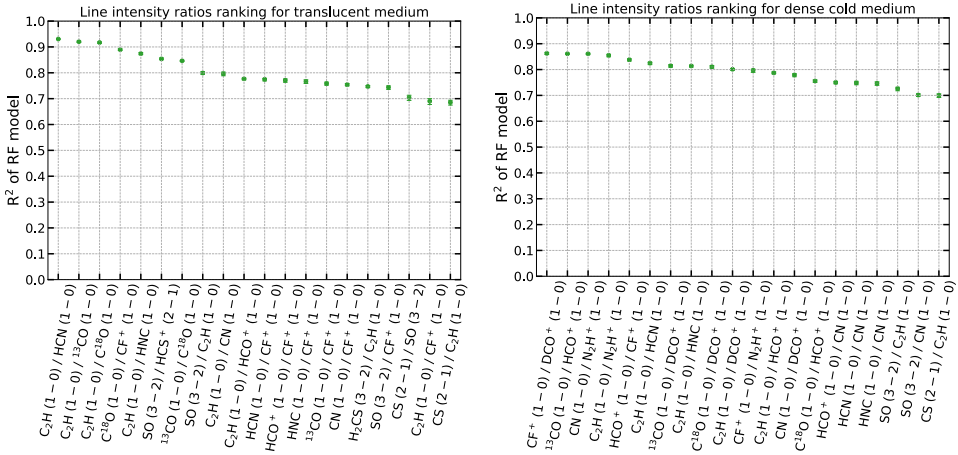
Massive datasets of model results produced by running large grids of models exploring wide ranges for multiple parameters and predicting tens or hundreds of observable quantities are now becoming common. These datasets contain a wealth of information on how sensitive the different observable quantities are to different unobservable physical quantities, and can thus help in finding new observable tracers of unobservable quantities such as  $x(e^-)$ .

The possible relationship between one observable quantity (e.g. an integrated line intensity ratio) and one unobservable quantity of interest (e.g. the ionization fraction) will however be affected by other unobservable parameters (e.g. the gas density, temperature, UV field, etc...). These will act as sources of noise and cause a scatter in the relationship. The predictive power of an observable quantity to predict a given unobservable quantity (here  $x(e^-)$ ) can thus be quantified as the magnitude of this scatter around the mean trend of the relationship. We now describe successively the two components needed for this purpose: a grid of models sampling randomly all of the important unobservable parameters over realistic ranges of values for the environment of interest, and a statistical method to find the mean trend of this relationship and to quantify the scatter around this mean trend.

**Table 1.** Ranges of parameters explored in the two model grids.

|                                     | Translucent medium grid                     | Cold dense medium grid                      |
|-------------------------------------|---|---|
| $n_{\text{H}}$ [ $\text{cm}^{-3}$ ] | $3 \times 10^2 - 3 \times 10^3$             | $10^3 - 10^6$                               |
| $T_{\text{gas}}$ [K]                | 15 – 100                                    | 7 – 20                                      |
| $G_0$                               | 1 – 1000                                    | 1   |
| $A_{\text{v}}$                      | 2 – 6                                       | 5 – 20                                      |
| $\zeta$ [ $\text{cm}^{-1}$ ]        | $10^{-17} - 10^{-15}$                       | $10^{-17} - 10^{-16}$                       |
| $\text{OPR}_{\text{H}_2}$           | 0.1 – 3                                     | $10^{-4} - 0.1$                             |
| depletion factor                    | 1   | 1 – 10                                      |
| [S]                                 | $1.86 \times 10^{-8} - 1.86 \times 10^{-5}$ | $1.86 \times 10^{-8} - 1.86 \times 10^{-5}$ |

Using the astrochemical model presented in [8], we compute single zone models with fixed density and temperature at stationary state, including a chemical network of 310 species and 8711 chemical reactions. As we expect the best tracers of the ionization fraction  $x(e^-)$  to be different in different types of environment, we run two grids of 5000 models each, exploring ranges of parameter values relevant respectively for translucent medium and for cold dense medium (the ranges of parameters explored are listed in Table 1). We select a list of species based on the frequency range of the Orion-B Large Program : CS, SO,  $\text{C}_2\text{H}$ ,



**Figure 1.**  $R^2$  value of the line intensity ratios as tracers of the ionization fraction in translucent medium conditions (left panel) and cold dense medium conditions (right panel).

<sup>13</sup>CO, C<sup>18</sup>O, HCO<sup>+</sup>, HCN, HNC, CN, H<sub>2</sub>CS, HCS<sup>+</sup>, CF<sup>+</sup> (and N<sub>2</sub>H<sup>+</sup> and DCO<sup>+</sup> for the cold dense medium conditions only). We post-process the results of each model using a non-LTE excitation and radiative transfer model (RADEX, [9]) to compute the integrated line intensity of these species and take as observable quantities all possible ratios between these integrated line intensities. More details on the grids of models used can be found in [1].

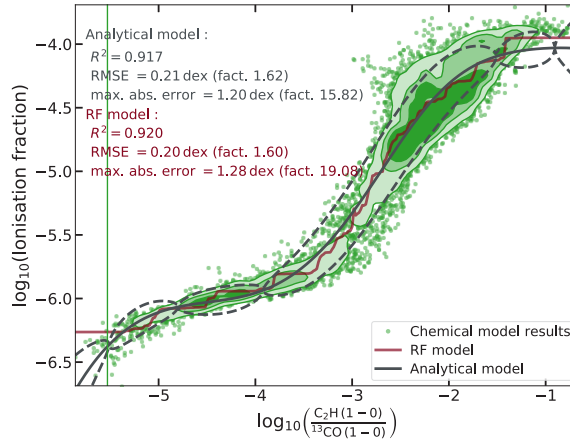
In order to capture the mean trend of the relationship between each line ratio and  $x(e^-)$ , we train a Random Forest statistical model [2] to learn to predict  $x(e^-)$  knowing only one given line ratio, based on the dataset of model results. To quantify the predictive power of the given line ratio, we use the  $R^2$  value ("out-of-bag" estimate) of the trained predictive model. For a predicted quantity  $y$ ,

$$R^2 = 1 - \frac{\sum_i (y_i^{\text{pred}} - y_i^{\text{true}})^2}{\sum_i (y_i^{\text{true}} - \overline{y^{\text{true}}})^2},$$

where the sums are over the datapoints, "pred" denotes the prediction while "true" the true value in the dataset, and  $\overline{y^{\text{true}}}$  is the mean of  $y$  in the dataset. The  $R^2$  value thus represents the fraction of the initial variance in the quantity of interest (here  $x(e^-)$ ) in the dataset that can be predicted from the fitted relationship with the single line ratio under consideration. We fit one Random Forest model and compute the corresponding  $R^2$  coefficient for each line ratio (separately for each of the two model grids), and then rank the tracers based on the  $R^2$  values. More details on how we trained the Random Forest models can be found in [1].

### 3 Results

The rankings of line ratios based on the  $R^2$  value are shown on Fig. 1 (left panel: translucent medium conditions, right panel: cold dense medium conditions). Several line ratios can predict a large fraction of the variations of  $x(e^-)$  in the model grids. In the translucent case, most of the best ratios involve C<sub>2</sub>H, which is known to be enhanced in UV-illuminated environments (e.g. [5, 6]). In translucent gas, UV photons are still the main source of ionization, and a tracer of the impact of the UV field such as C<sub>2</sub>H can thus also trace accurately the ionization



**Figure 2.** Scatterplot of the ionization fraction versus the  $C_2H(1-0)/^{13}CO(1-0)$  line intensity ratio for the translucent medium grid of models. Superimposed is the mean trend captured by the Random Forest (red line), with measures of the goodness of fit ( $R^2$ , mean squared error and maximum absolute error given on the figure). A simpler analytical model described in [1] and fitted to the data is also superimposed (black lines). Credit : Bron, A&A, 645, A28, 2021, reproduced with permission ©ESO.

fraction of the gas. In both types of conditions, we find very good tracers of the ionization fraction, that can give accurate estimates of the ionization fraction despite not knowing the density, temperature, UV illumination, ortho-para ratio of  $H_2$ , etc. For instance, Fig. 2 shows a scatter plot of the ionization fraction versus one of the best tracers in the translucent case: the  $C_2H(1-0)/^{13}CO(1-0)$  line ratio. The scatter induced by the 7 unknown parameters in the model grid remains very limited, as can be expected from the  $R^2 > 0.9$  found. An ad hoc analytical fit, more easy to reuse than the Random Forest model, is also provided (black solid line, uncertainties as dashed lines). Complete results can be found in [1].

The method that we propose is very general and can be applied to find good tracers of any physical quantities (e.g., gas density, cosmic ray ionization rate) based on datasets of any types of more complex models (e.g., PDR models, astrochemical models including surface chemistry). A general implementation of this method is publicly available at <http://autorank.ism.obspm.fr/>.

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