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Magnetic topologies of young suns: the weak-line T Tauri stars TWA 6 and TWA 8A

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ABSTRACT

We present a spectropolarimetric study of two weak-line T Tauri stars (wTTs), TWA 6, and TWA 8A, as part of the MaTYSSE (Magnetic Topologies of Young Stars and the Survival of close-in giant Exoplanets) programme. Both stars display significant Zeeman signatures that we have modelled using Zeeman Doppler Imaging (ZDI). The magnetic field of TWA 6 is split equally between poloidal and toroidal components, with the largest fraction of energy in higher order modes, with a total unsigned flux of 840 G, and a poloidal component tilted 35° from the rotation axis. TWA 8A has a 70 per cent poloidal field, with most of the energy in higher order modes, with an unsigned flux of 1.4 kG (with a magnetic filling factor of 0.2), and a poloidal field tilted 20° from the rotation axis. Spectral fitting of the very strong field in TWA 8A (in individual lines, simultaneously for Stokes *I* and *V*) yielded a mean magnetic field strength of 5.9 ± 0.2 kG. The higher field strengths recovered from spectral fitting suggests that a significant proportion of magnetic energy lies in small-scale fields that are unresolved by ZDI. So far, wTTs in MaTYSSE appear to show that the poloidal-field axisymmetry correlates with the magnetic field strength. Moreover, it appears that classical T Tauri stars (cTTs) and wTTs are mostly poloidal and axisymmetric when mostly convective and cooler than ~ 4300 K, with hotter stars being less axisymmetric and poloidal, regardless of internal structure.

Key words: techniques: polarimetric – stars: formation – stars: imaging – stars: individual: TWA 6 – stars: individual: TWA 8A – magnetic fields.

1 INTRODUCTION

During the first few hundred thousand years of low-mass star formation, class-I pre-main sequence (PMS) stars accrete significant amounts of material from their surrounding dusty envelopes. After around 0.5 Myr, these protostars emerge from their dusty cocoons and are termed classical T Tauri stars (cTTs/class-II PMS stars) if they are still accreting from their surrounding discs, or weak-line T Tauri stars (wTTs/class-III PMS stars) if they have exhausted the gas from the inner disc cavity. During the PMS phase, stellar magnetic fields have their largest impact on the evolution of the star. These fields control accretion processes and trigger outflows/jets (Bouvier et al. 2007), dictate the star’s angular momentum evolution

by enforced spin-down through star-disc coupling (e.g. Davies, Gregory & Greaves 2014), and alter disc dynamics and planet formation (Baruteau et al. 2014). Moreover, as PMS stars are gravitationally contracting towards the MS, the change in stellar structure from fully to partly convective is expected to alter the stellar dynamo mechanism and the resulting magnetic field topology.

Previous work through the MaPP (Magnetic Protostars and Planets) survey revealed that the large-scale topologies of 11 cTTs remained relatively simple and mainly poloidal when the host star is still fully or largely convective, but become much more complex when the host star turns mostly radiative (Gregory et al. 2012; Donati et al. 2013). This survey concluded that these fields likely originated from a dynamo, varying over time-scales of a few years (Donati et al. 2011, 2012, 2013), and resembling

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those of mature stars with comparable internal structure (Morin et al. 2008).

The nature of the magnetic fields of wTTSs and how they depend on fundamental parameters is less well known. These evolutionary phases are the initial conditions in which discless PMS stars initiate their unleashed spin-up towards the zero-age main sequence (ZAMS). Hence, it is crucial to characterize their magnetic fields and how they depend on mass, temperature, age, and rotation. To this end, we are performing a spectropolarimetric study of around 30 wTTSs through the MaTYSSSE (Magnetic Topologies of Young Stars and the Survival of close-in giant Exoplanets) programme, mainly allocated on ESPaDOnS at the Canada–France–Hawaii Telescope (CFHT), complemented by observations with NARVAL on the Telescope Bernard Lyot, and with HARPS on the ESO 3.6-m Telescope. By using Zeeman Doppler Imaging (ZDI) to characterize the magnetic fields of wTTSs, we are able to test stellar dynamo theories and models of low-mass star formation. Moreover, by filtering out the activity-related jitter from the radial velocity (RV) curves, we are able to potentially detect hot Jupiters (hJs; see Donati et al. 2016), and thus verify whether core accretion and migration is the most likely mechanism for forming close-in giant planets (e.g. Alibert et al. 2005).

Here, we present our detailed analysis of the wTTSs TWA 6 and TWA 8A as part of the MaTYSSSE survey. Both targets are members of the TW Hydrae association, which, at an age of 10 ± 3 Myr (Bell, Mamajek & Naylor 2015), is in transition between the T Tauri and the post T Tauri phase, and thus provides a very interesting period in which to study the properties of the member stars as they spin-up towards the ZAMS. Our phase-resolved spectropolarimetric observations are documented in Section 2, with the stellar and disc properties presented in Section 3. We discuss the spectral energy distributions (SEDs), several emission lines, and the accretion status of both stars in Section 3.2. In Section 4, we present our results after applying our tomographic modelling technique to the data. In Section 5, we present our results of our spectral fitting to the Stokes I and V spectra, and in Section 6, we discuss our analysis of the filtered RV curves. Finally, we discuss and summarize our results and their implications for low-mass star and planet formation in Section 7.

2 OBSERVATIONS

Spectropolarimetric observations of TWA 6 were taken in February 2014, with observations of TWA 8A taken in March and April 2015, both using ESPaDOnS at the 3.6-m CFHT. Spectra from ESPaDOnS span the entire optical domain (from 370 to 1000 nm) at a resolution of 65 000 (i.e. a resolved velocity element of 4.6 km s^{-1}) over the full wavelength range, in both circular or linear polarization (Donati 2003).

A total of 22 circularly polarized (Stokes V) and unpolarized (Stokes I) spectra were collected for TWA 6 over a time span of 16 nights, corresponding to around 29.6 rotation cycles (where $P_{\text{rot}} = 0.5409$ d; Kiraga 2012). Time sampling was fairly regular, with the longest gap of six nights occurring towards the end of the run. For TWA 8A, 15 spectra were collected with regular time sampling over a 15 night time span, corresponding to around 3.2 rotation cycles (where $P_{\text{rot}} = 4.638$ d; Kiraga 2012).

All polarization spectra consist of four individual sub-exposures (each lasting 406 s for TWA 6 and 1115 s for TWA 8A), taken in different polarimeter configurations to allow the removal of all spurious polarization signatures at first order. All raw frames

were processed using the LIBRE ESPRIT software package, which performs bias subtraction, flat fielding, wavelength calibration, and optimal extraction of (un)polarized Echelle spectra, as described in the previous papers of the series (Donati et al. 1997, also see Donati et al. 2010, 2011, 2014), to which the reader is referred for more information. The peak signal-to-noise ratios (S/N; per 2.6 km s^{-1} velocity bin) achieved on the collected spectra range between 111–197 (median 164) for TWA 6, and 209–369 (median 340) for TWA 8A, depending on weather/seeing conditions. All spectra are automatically corrected for spectral shifts resulting from instrumental effects (e.g. mechanical flexures, temperature or pressure variations) using atmospheric telluric lines as a reference. This procedure provides spectra with a relative RV precision of better than 0.030 km s^{-1} (e.g. Moutou et al. 2007; Donati et al. 2008). A journal of all observations is presented in Table 1 for both stars.

3 STELLAR AND DISC PROPERTIES

Both stars are part of the TW Hya association (TWA; e.g. Jayawardhana et al. 1999; Webb et al. 1999; Donaldson et al. 2016), one of the closest young star associations at a distance of $\simeq 50$ pc (e.g. Zuckerman & Song 2004). Furthermore, at an age of 10 ± 3 Myr (Bell et al. 2015), TWA is at a crucial evolutionary phase where star–disc interactions have ceased, and where the TTSs are rapidly spinning up as they continue their gravitational contraction towards the main sequence (e.g. Rebull, Wolff & Strom 2004).

Both stars are classed as T Tauri due to strong Li I 6708 Å absorption (e.g. Webb et al. 1999), with mean equivalent widths (EWs) of around 0.45 Å (20 km s^{-1}) and 0.38 Å (17 km s^{-1}) for TWA 6 and TWA 8A, respectively (slightly lower than the 0.56 and 0.53 Å found by Torres et al. 2003). Furthermore, our spectra show that the strength of Li I 6708 Å absorption does not vary significantly for either star, indicating a lack of veiling (in agreement with Herczeg & Hillenbrand 2014), and confirming their status as wTTSs (see Sections 3.2 and B for further discussion). Moreover, both stars show very regular periodic light curves that do not appear like those of cTTS, further supporting their non-accreting status.

For TWA 6, we adopt the photometric rotation period of 0.5409 d found by Kiraga (2012) for the remainder of this work, as this is in excellent agreement with the 0.54 ± 0.01 d period of Lawson & Crause (2005), and the 0.54090 ± 0.00005 d period of Skelly et al. (2008). For TWA 8A, we adopt the photometric period of 4.638 d (Kiraga 2012), in excellent agreement with the 4.65 ± 0.01 d period found by Lawson & Crause (2005), the 4.66 ± 0.06 d period of Messina et al. (2010), and the 4.639 d period found by applying a Lomb–Scargle periodogram analysis to SuperWASP photometric data (Butters et al. 2010). The rotational cycles of TWA 6 and TWA 8A (denoted E_1 and E_2 in equation 1) are computed from Barycentric Julian Dates (BJDs) according to the (arbitrary) ephemerides:

$$\begin{aligned} \text{BJD (d)} &= 2456693.9 + 0.5409E_1 \quad (\text{for TWA 6}), \\ \text{BJD (d)} &= 2457107.9 + 4.638E_2 \quad (\text{for TWA 8A}). \end{aligned} \quad (1)$$

3.1 Stellar properties

To determine the T_{eff} and $\log g$ of our target stars, we applied our automatic spectral classification tool (discussed in Donati et al. 2012) to several of the highest S/N spectra for both stars.

Table 1. Journal of ESPaDOnS observations of TWA 6 (first 22 rows) and TWA 8A (last 15 rows), each consisting of a sequence of four subexposures lasting 406 s and 1115 s for TWA 6 and TWA 8A, respectively. Columns 1–4 list (i) the UT date of the observation, (ii) the corresponding UT time at mid-exposure, (iii) the Barycentric Julian Date (BJD), and (iv) the peak S/N (per 2.6 km s⁻¹ velocity bin) of each observation. Columns 5 and 6, respectively, list the S/N in Stokes/LSD profiles (per 1.8 km s⁻¹ velocity bin), and the rms noise level (relative to the unpolarized continuum level I_c and per 1.8 km s⁻¹ velocity bin) in the Stokes V LSD profiles. Column 6 indicates the rotational cycle associated with each exposure, using the ephemerides given in equation (1).

Date (2014)	UT (hh:mm:ss)	BJD (2456693.9+)	S/N	S/N _{LSD}	σ_{LSD} (0.01 per cent)	Cycle
Feb 04	11:15:00	0.07239	160	1796	5.6	0.151
Feb 04	12:17:52	0.11605	184	2260	4.5	0.232
Feb 07	10:28:34	3.04027	134	1619	6.2	5.638
Feb 07	11:30:04	3.08299	131	1559	6.4	5.717
Feb 07	13:04:25	3.14851	158	1849	5.4	5.838
Feb 09	09:27:53	4.99822	168	1990	5.1	9.258
Feb 09	10:46:58	5.05313	169	1973	5.1	9.359
Feb 09	11:48:45	5.09605	132	1639	6.1	9.439
Feb 10	11:08:25	6.06807	163	1914	5.3	11.236
Feb 10	12:10:38	6.11128	178	2126	4.7	11.316
Feb 11	09:37:38	7.00507	183	2273	4.4	12.968
Feb 11	11:05:43	7.06624	197	2461	4.1	13.081
Feb 11	11:54:04	7.09981	169	1922	5.2	13.143
Feb 12	09:21:45	7.99407	164	1933	5.2	14.797
Feb 12	11:22:19	8.07780	159	1869	5.4	14.951
Feb 12	12:49:30	8.13834	160	1855	5.4	15.063
Feb 13	10:35:31	9.04533	111	1514	6.6	16.740
Feb 13	13:06:11	9.14997	180	2094	4.8	16.934
Feb 19	09:23:25	14.99545	160	1847	5.4	27.741
Feb 19	10:51:16	15.05645	181	2223	4.5	27.853
Feb 19	12:29:16	15.12451	192	2387	4.2	27.979
Feb 20	12:06:25	16.10867	137	1345	7.5	29.799
(2015)		(2457107.9+)				
Mar 25	11:43:05	0.06756	338	3847	2.6	0.020
Mar 26	11:04:43	1.04090	343	3812	2.6	0.230
Mar 27	11:40:03	2.06545	340	3863	2.6	0.451
Mar 28	11:30:18	3.05868	341	3841	2.6	0.665
Mar 29	12:16:00	4.09040	302	3348	3.0	0.887
Mar 30	11:49:48	5.07220	369	4244	2.4	1.099
Mar 31	08:28:33	5.93245	357	4071	2.5	1.285
Apr 01	08:26:55	6.93130	349	3960	2.5	1.500
Apr 03	11:25:50	9.05552	253	2670	3.8	1.958
Apr 04	11:34:15	10.06136	355	4091	2.5	2.175
Apr 05	08:42:10	10.94184	332	3764	2.7	2.365
Apr 06	08:30:36	11.93379	353	4013	2.5	2.579
Apr 08	09:13:10	13.97061	202	1754	5.7	3.018
Apr 09	07:18:42	14.90036	253	2819	3.6	3.218
Apr 09	08:45:35	14.95143	309	3621	2.8	3.229

We fit the observed spectrum using multiple windows in the wavelength ranges 515–520 and 600–620 nm (using Kurucz model atmospheres; Kurucz 1993), in a similar way to the method of Valenti & Fischer (2005). This process yields estimates of T_{eff} and $\log g$, where the optimum parameters are those that minimize χ^2 , with errors bars determined from the curvature of the χ^2 landscape at the derived minimum.

For TWA 6, we find that $T_{\text{eff}} = 4425 \pm 50$ K and $\log g = 4.5 \pm 0.2$ (with g in cgs units). While two-temperature modelling such as that carried out by Gully-Santiago et al. (2017) would provide a better estimate of T_{eff} and the fractional spot coverage, for our purposes, a homogeneous model is sufficient. For TWA 6, we adopt the V and B magnitudes of 10.88 ± 0.05 and 12.19 ± 0.05 from (Messina et al. 2010), and assuming a spot coverage of the visible stellar hemisphere of ~ 30 per cent (typical for such

active stars, see Section 4), we derive an unspotted V magnitude of 10.6 ± 0.2 . We note that assuming a different spot coverage (such as 0 or 50 per cent) places our derived parameters within our quoted error bars. Using the relation from Pecaut & Mamajek (2013), the expected visual bolometric correction for TWA 6 is $BC_V = -0.70 \pm 0.04$, and as there is no evidence of extinction to TWA members (e.g. Stelzer et al. 2013), we adopt $A_V = 0$. Combining V , BC_V , A_V and the trigonometric parallax distance found by *Gaia* of 63.9 ± 1.4 pc (corresponding to a distance modulus of 4.03 ± 0.05 , *Gaia* Collaboration et al. 2016, 2018, in excellent agreement with the 59.59 ± 3.6 pc of Donaldson et al. 2016), we obtain an absolute bolometric magnitude of 5.85 ± 0.29 , or equivalently, a logarithmic luminosity relative to the Sun of -0.44 ± 0.12 . When combined with the photospheric temperature obtained previously, we obtain a radius of $1.0 \pm 0.2 R_{\odot}$.

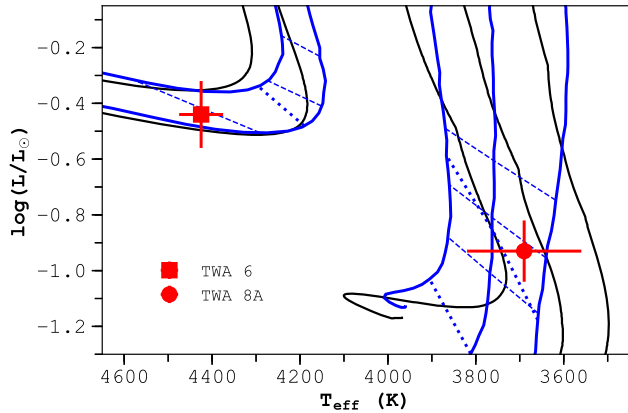


Figure 1. H–R diagram showing the stellar evolutionary tracks provided by Siess et al. (2000, blue solid lines) and Baraffe et al. (2015, black solid lines) for masses of 0.3, 0.4, 0.5, 0.9, and 1.0 M_{\odot} . Blue dashed lines show the corresponding isochrones for ages 5, 10, and 20 Myr, and blue dotted lines mark the 0 and 50 per cent fractional radius for the bottom of the convective envelope, both for Siess et al. (2000) models.

Coupling P_{rot} (see equation 1) with the measured $v \sin i$ of $72.6 \pm 0.5 \text{ km s}^{-1}$ (see Section 4), we can infer that $R_* \sin i$ is equal to $0.78 \pm 0.01 R_{\odot}$, where R_* and i denote the stellar radius and the inclination of its rotation axis to the line of sight. By comparing the luminosity-derived radius to that from the stellar rotation, we derive that i is equal to 49_{-8}^{+15} , in excellent agreement with that found using our tomographic modelling (see Section 4). Using the evolutionary models of Siess, Dufour & Forestini (2000) (assuming solar metallicity and including convective overshooting), we find that TWA 6 has a mass of $0.95 \pm 0.10 M_{\odot}$, with an age of 21 ± 9 Myr (see the H–R diagram in Fig. 1, with evolutionary tracks and corresponding isochrones). Similarly, using the evolutionary models of Baraffe et al. (2015), we obtain a mass of $0.95 \pm 0.10 M_{\odot}$ and an age of 17 ± 7 Myr.

For TWA 8A, our spectral fitting code yields a best fit at $T_{\text{eff}} = 3800 \pm 150 \text{ K}$ and $\log g = 4.7 \pm 0.2$; however, this T_{eff} is in the regime where the Kurucz synthetic spectra are considered unreliable in terms of temperature. To address this issue, we are currently working on a more advanced spectral classification tool based on PHOENIX model atmospheres and synthetic spectra (see Allard 2014). In the mean time for the work presented here, we determined T_{eff} for TWA 8A from the observed $B - V$ value and the relation between T_{eff} and $B - V$ for young stars from Pecauc & Mamajek (2013) (and by assuming $A_V = 0$). We adopt $V = 12.265 \pm 0.023$ and $B = 13.70 \pm 0.03$ from Henden et al. (2015), with $B - V = 1.434 \pm 0.038$. Using this $B - V$ with the relation between intrinsic colour and T_{eff} for young stars found by Pecauc & Mamajek (2013), and assuming $A_V = 0$, we derive $T_{\text{eff}} = 3690 \pm 130 \text{ K}$. Combining the observed V magnitude with the expected BC_V for TWA 8A of -1.50 ± 0.19 (Pecauc & Mamajek 2013) with the trigonometric parallax distance of $46.27 \pm 0.19 \text{ pc}$ as found by Gaia (Gaia Collaboration et al. 2016, 2018, corresponding to a distance modulus of 3.326 ± 0.009 , in excellent agreement with the $47.2 \pm 2.8 \text{ pc}$ of Donaldson et al. 2016 and $46.9_{-2.9}^{+3.3} \text{ pc}$ of Riedel et al. 2014), we obtain an absolute bolometric magnitude of 7.1 ± 0.3 , or equivalently, a logarithmic luminosity relative to the Sun of -0.93 ± 0.11 . When combined with the photospheric temperature obtained previously, we obtain a radius of $0.8 \pm 0.2 R_{\odot}$. Combining this radius with the mass derived below (from Baraffe et al. 2015 evolutionary models), we estimate $\log g = 4.3 \pm 0.3$.

Combining P_{rot} (see equation 1) with the $v \sin i$ of $4.82 \pm 0.16 \text{ km s}^{-1}$ (see Section 5), we find $R_* \sin i = 0.44 \pm 0.03 R_{\odot}$, yielding $i = 32_{-8}^{+13}$, in good agreement with our tomographic modelling (see Section 4). Using Siess et al. (2000) models we find $M = 0.45 \pm 0.10 M_{\odot}$, with an age of 11 ± 5 Myr. Using the evolutionary models of Baraffe et al. (2015), we find $M = 0.55 \pm 0.1 M_{\odot}$, with an age of 13 ± 6 Myr.

We note that we do not consider the formal error bars on the derived masses and ages to be representative of the true uncertainties, given the inherent limitations of these evolutionary models. Furthermore, we note that for internal consistency with previous MaPP and MaTYSSSE results, the values from the Siess et al. (2000) models should be referenced. We note that the ages derived here are consistent with the age of the young TWA moving group (of 10 ± 3 Myr; Bell et al. 2015), and that both evolutionary models suggest that TWA 6 has a mostly radiative interior, whereas TWA 8A is mostly (or fully) convective.

The temperatures measured here are hotter than expected from spectral types estimated from red-optical spectra that cover TiO and other molecular bands (White & Hillenbrand 2004; Stelzer et al. 2013; Herczeg & Hillenbrand 2014). This discrepancy is consistent with past wavelength-dependent differences in photospheric temperatures from young stars, which may be introduced by spots (e.g. Bouvier & Appenzeller 1992; Debes et al. 2013; Gully-Santiago et al. 2017). The interpretation of these differences is not yet understood. Use of the lower temperatures that are measured at longer wavelengths from molecular bands would lead to lower masses and younger ages. Our temperatures are accurate measurements of the photospheric emission from 5000 to 6000 Å and are consistent with all temperature measurements for stars in the MaTYSSSE programme.

3.2 Spectral energy distributions

SEDs of TWA 6 and TWA 8A were constructed using photometry sourced from the DENIS survey (G95), the AAVSO Photometric All Sky Survey (APASS; Henden et al. 2015), the GALEX all-sky imaging survey (Bianchi et al. 2011), the TYCHO-2 catalogue (Høg et al. 2000), the WISE, Spitzer and Gaia catalogues (Wright et al. 2010; Werner et al. 2004; Gaia Collaboration et al. 2016, 2018), and Torres et al. (2006). We note that deep, sensitive sub-mm and mm photometry are not currently available for our targets. Comparing the SEDs (shown in Fig. 2) to PHOENIX-BT-Settl synthetic spectra (Allard 2014), we find that neither TWA 6 nor TWA 8A has an infrared (IR) excess up to $23.675 \mu\text{m}$, indicating that both objects have dissipated their circumstellar discs. Given that the SEDs of TWA 6 and TWA 8A show no evidence of an IR excess, both stars are likely discless and are not accreting (also see e.g. Weinberger et al. 2004; Low et al. 2005). However, for completeness, in Appendix B we present several metrics that determine the accretion rates from emission lines (if accretion were present), with our analysis showing that chromospheric emission likely dominates the line formation for both targets, confirming their classification as WTTSs.

3.3 Emission-line analysis

We find that TWA 6 shows core Ca II infrared triplet (IRT) emission (see Fig. A1) with a mean EW of around 0.3 \AA (10.7 km s^{-1}), similar to what is expected from chromospheric emission for such PMS stars (e.g. Ingleby et al. 2011), and lower than that for accreting cTTs (e.g. Donati et al. 2007). The core Ca II IRT emission is somewhat variable, with both red and blue-shifted peaks (where the

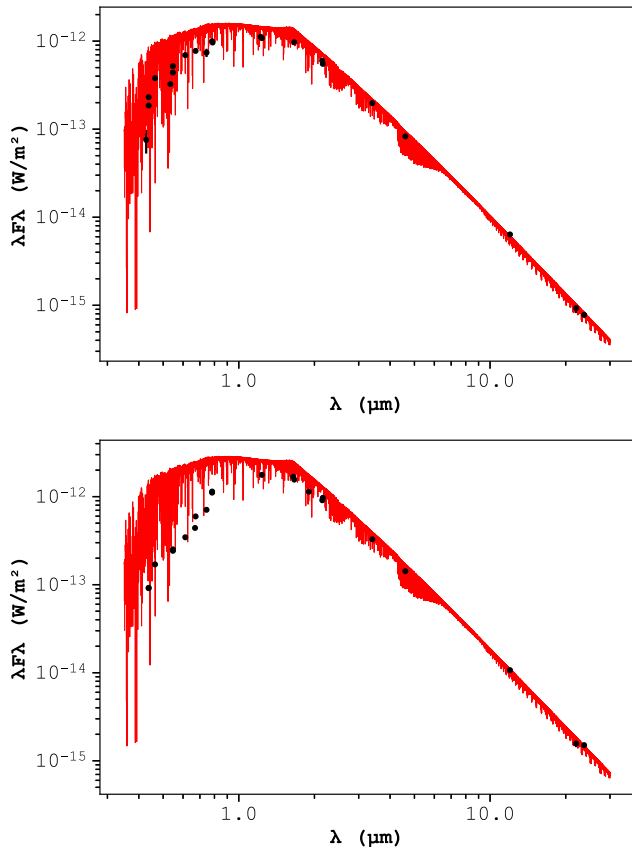


Figure 2. SEDs of TWA 6 (top) and TWA 8A (bottom), where the photometric data (see the text) are shown as black dots, and where PHOENIX-BT-Settl model spectra (Allard 2014) are shown as a red line. For the model spectra, we adopt $T_{\text{eff}} = 4400$ and 3700 K for TWA 6 and TWA 8A, respectively, and $\log g = 4.5$ for both stars, as well as the other parameters given in Table 2, adopting the extinction relation of Cardelli, Clayton & Mathis (1989). Furthermore, we assume that both stars have a 30 percent surface coverage of cool starspots (see Section 4.1), and so the displayed spectra have a 30 per cent contribution from a spectrum that is 1000 K cooler.

red-shifted emission is generally larger), and where the emission is significantly higher at cycles 9.258, 9.359, 14.951, and 15.063. We note that there are some differences in the Stokes V line profiles of the Ca II IRT that are likely due to their different atmospheric formation heights. We note that no significant Zeeman signatures are detected in Ca II H&K, Ca II IRT, or He I 5875.62 Å and so the emission is likely chromospheric rather than from the magnetic footpoints of an accretion funnel. TWA 6 also shows single-peaked H α and H β emission that displays relatively little variability over the ~ 30 rotation cycles (see Fig. A1). For H α , significantly higher flux is seen in cycles 9.258, 9.359, and 9.439, with the extra emission arising in a predominantly red-shifted component. Moreover, cycle 14.797 displays a significantly higher flux that is symmetric about zero velocity. This higher flux is also seen in H β , with larger emission for cycles 9.258 and 9.359 (both asymmetric and red-shifted), 14.797 (symmetric), and 14.951 (asymmetric, red-shifted). Given that these emission features occur at similar phases in Ca II IRT, H α , and H β , and are also short lived, they likely stem from the same formation mechanism in the form of stellar prominences that are rotating away from the observer. This conclusion is also supported by the mapped magnetic topology, as we see closed magnetic loops off the stellar limb, along which

prominence material may flow. To better determine the nature of the emission and its variability, one can calculate variance profiles and autocorrelation matrices, as described in Johns & Basri (1995) and given by

$$V_{\lambda} = \left[\frac{\sum_{i=1}^n (I_{\lambda,i} - \bar{I}_{\lambda})^2}{n-1} \right]^{\frac{1}{2}}. \quad (2)$$

Fig. A4 shows that the H α emission varies from around -200 to $+300$ km s $^{-1}$ (similar to that found previously for TWA 6 by Skelly et al. 2008), well beyond the $v \sin i$ of 72.6 km s $^{-1}$, and with most of the variability in a red-shifted component. Furthermore, the autocorrelation matrix shows strong correlation of the low-velocity components, indicating a common origin. We find that H β and He I D3 show negligible variability, with a relatively low-spectral S/N limiting the analysis.

In the case of TWA 8A, core Ca II IRT emission is present with a mean EW of around 0.37 Å (11.9 km s $^{-1}$, see Fig. A2). This emission is mostly non-variable, with only cycle 1.958 showing significantly higher (symmetric) emission. Furthermore, the Zeeman signatures in the Stokes V line profiles (see Fig. A3) have the same sign as those of the absorption lines (see Fig. 3), and so are of photospheric origin. TWA 8A also displays double-peaked H α and H β emissions, with a peak separation of around 40 km s $^{-1}$. This separation lies well within the co-rotation radius, and is only a few times larger than the $v \sin i$ of 4.82 km s $^{-1}$, indicating that the source of the emission is chromospheric. The lines are somewhat variable, with a significant increase in emission (for both H α and H β) at cycles 1.958, 2.579, and 3.018. Fig. A5 shows the variance profiles and autocorrelation matrices of H α , H β , and He I D3. Here, we see that for H α , the variability concentrates in two peaks centred around -50 and $+75$ km s $^{-1}$ (ranging ± 150 km s $^{-1}$), with variability in H β likewise occurring in two peaks centred around -75 and $+65$ km s $^{-1}$ (ranging ± 150 km s $^{-1}$), with both autocorrelation matrices showing the low-velocity components to be highly correlated. For He I D3, we find that the variability is single peaked, centred around zero velocity, with only low-velocity components showing significant correlation. We also note that the H α emission of TWA 8A shows strong Zeeman signatures (see Fig. A3) that are opposite in sign to those of the absorption lines (see Fig. 5), as expected for chromospheric emission.

4 TOMOGRAPHIC MODELLING

In order to map, both the surface brightness and magnetic field topology of TWA 6 and TWA 8A, we have applied our dedicated stellar-surface tomographic-imaging package to the data sets described in Section 2. In doing this, we assumed that the observed variability is dominated by rotational modulation (and optionally differential rotation). Our imaging code simultaneously inverts the time series of Stokes I and V profiles into brightness maps (featuring both cool spots and warm plages) and magnetic maps (with poloidal and toroidal components, using a spherical harmonic decomposition). For brightness imaging, a copy of a local line profile is assigned to each pixel on a spherical grid, and the total line profile is found by summing over all visible pixels (at a given phase), where the pixel intensities are scaled iteratively to fit the observed data. For magnetic imaging, the Zeeman signatures are fit using a spherical-harmonic (SH) decomposition of potential and toroidal field components, where the weighting of the harmonics are scaled iteratively (Donati 2001). The data are fit to an aim χ^2 , with the optimal fit determined using the maximum-entropy routine

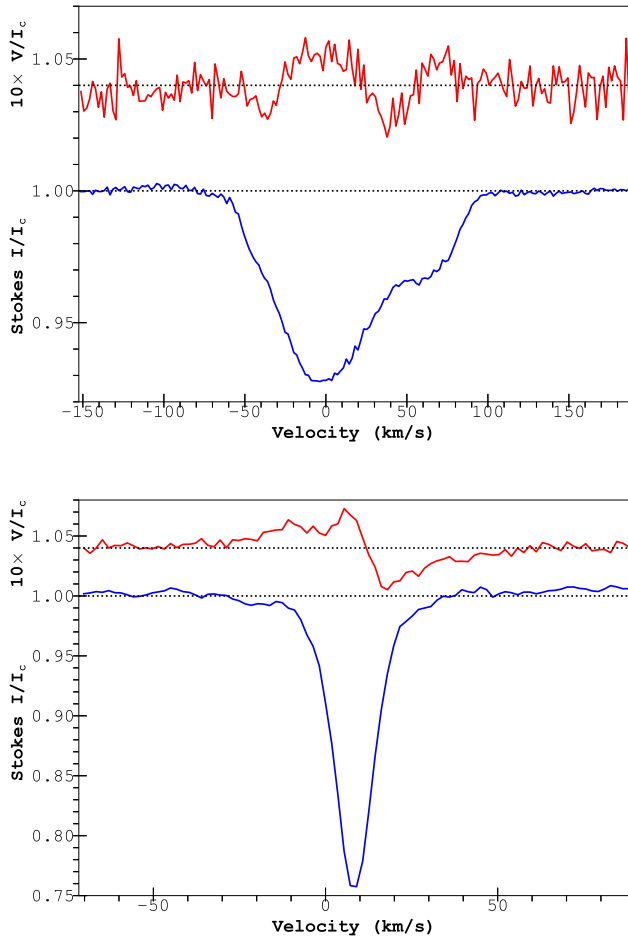


Figure 3. LSD circularly polarized (Stokes V , top/red curve) and unpolarized (Stokes I , bottom/blue curve) profiles of TWA 6 (top, collected on 19-02-2014, cycle 27.979) and TWA 8A (bottom, collected on 26-03-2015, cycle 0.223). Clear Zeeman signatures are detected in both LSD Stokes V profiles in conjunction with the unpolarized line profiles. The mean polarization profiles are expanded by a factor of 10 shifted upwards by 0.04 for display purposes.

of Skilling & Bryan (1984), and where the chosen map is that which contains least information (where entropy is maximized) required to fit the data. For further details about the specific application of our code to wTTSs, we refer the reader to previous papers in the series (e.g. Donati et al. 2010, 2014, 2015).

As with previous studies of wTTSs, we applied the technique of least-squares deconvolution (LSD; Donati et al. 1997) to all of our spectra. Given that relative noise levels are around 10^{-3} in a typical spectrum (for a single line), with Zeeman signatures exhibiting relative amplitudes of ~ 0.1 per cent, the use of LSD allows us to create a single ‘mean’ line profile with a dramatically enhanced S/N, with accurate error bars for the Zeeman signatures. LSD involves cross-correlating the observed spectrum with a stellar line list, and for this work, stellar line lists were sourced from the Vienna Atomic Line Database (VALD; Ryabchikova et al. 2015), computed for $T_{\text{eff}} = 4500$ K and $\log g = 4.5$ (in cgs units) for TWA 6, and $T_{\text{eff}} = 3750$ K and $\log g = 4.5$ for TWA 8A (the closest available to our derived spectral types; see Section 3.1). Only moderate-to-strong atomic spectral lines were included (with line-to-continuum core depressions larger than 40 per cent prior to all non-thermal broadening). Furthermore, spectral regions containing

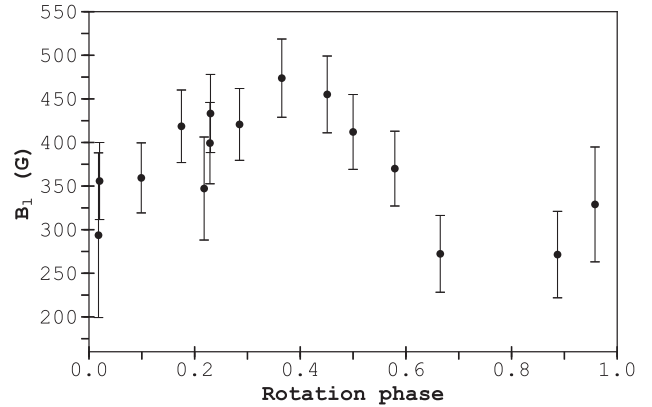


Figure 4. The longitudinal field strengths $\langle B_l \rangle$ for TWA 8A, as measured from the LSD profiles.

strong lines mostly formed outside the photosphere (e.g. Balmer, He, Ca II H&K, and Ca II IRT lines) and regions heavily crowded with telluric lines were discarded (see e.g. Donati et al. 2010 for more details), leaving 6088 and 5953 spectral lines for use in LSD, for TWA 6 and TWA 8A, respectively. Expressed in units of the unpolarized continuum level I_c (and per 1.8 km s^{-1} velocity bin), the average noise level of the resulting Stokes V signatures range from 4.1 to 7.5×10^{-4} (median of 5.3×10^{-4}) for TWA 6 and 2.4 to 5.7×10^{-4} (median of 2.6×10^{-4}) for TWA 8A.

The disc-integrated average photospheric LSD profiles are computed by first synthesizing the local Stokes I and V profiles using the Unno-Rachkovsky analytical solution to the polarized radiative transfer equations in a Milne-Eddington model atmosphere, taking into account the local brightness and magnetic field. Then, these local line profiles are integrated over the visible hemisphere (including linear limb darkening, with a coefficient of 0.75, as observed young stars, e.g. Donati & Collier Cameron 1997) to produce synthetic profiles for comparison with observations. This method provides a reliable description of how line profiles are distorted due to magnetic fields (including magneto-optical effects, e.g. Landi Degl’Innocenti & Landolfi 2004). The main parameters of the local line profiles are similar to those used in our previous studies; the wavelength, Doppler width, EW, and Landé factor being set to 670 nm, 1.8 km s^{-1} , 3.9 km s^{-1} , and 1.2, respectively.

We note that while Zeeman signatures are detected at all times in Stokes V LSD profiles for both stars (see Fig. 3 for an example), TWA 8A exhibits much larger longitudinal field strengths (B_l), similar to those of e.g. mid-M dwarfs (see Morin et al. 2008), with values shown in Fig. 4, as calculated from the LSD profiles. Here, we clearly see the periodicity in field strength, with the maximum B_l around phase 0.37, coincident with the phase of the aligned dipole of the magnetic field (see Fig. 7) being viewed along the line of sight, with the minimum B_l seen around half a rotation later. TWA 8A also exhibits significant Zeeman broadening in the Stokes I profiles that we model in Section 5, with almost no distortions due to brightness inhomogeneities on the surface.

As part of the imaging process, we obtain accurate estimates for v_{rad} (the RV the star would have if unspotted), equal to 17.5 ± 0.1 and $8.34 \pm 0.10 \text{ km s}^{-1}$, the inclination i of the rotation axis to the line of sight, equal to $46^\circ \pm 10^\circ$ and $42^\circ \pm 10^\circ$, for TWA 6 and TWA 8A, respectively, and for TWA 6 the $v \sin i$ equal to $72.6 \pm 0.1 \text{ km s}^{-1}$ (see Table 2, in excellent agreement with the

Table 2. Main parameters of TWA 6 and TWA 8A as derived from our study, with v_{rad} noting the RV that the star would have if unspotted, the equatorial rotation rate Ω_{eq} and the difference between equatorial and polar rotation rates $d\Omega$ (as inferred from the modelling of Section 4). Note, the stellar masses and ages are those determined from Siess et al. (2000) models, with values from Baraffe et al. (2015) given in parenthesis. The $\log g$ for TWA 8A is estimated from its mass [using Baraffe et al. (2015) models] and R_* .

	TWA 6	TWA 8A
M_* (M_{\odot})	0.95 ± 0.10 (0.95 \pm 0.10)	0.45 ± 0.10 (0.55 \pm 0.10)
R_* (R_{\odot})	1.1 ± 0.2	0.8 ± 0.2
Age (Myr)	21 ± 9 (17 \pm 7)	11 ± 5 (13 \pm 6)
$\log g$ (cgs units)	4.5 ± 0.2	4.3 ± 0.3
T_{eff} (K)	4425 ± 50	3690 ± 130
$\log(L_*/L_{\odot})$	-0.40 ± 0.12	-0.93 ± 0.11
P_{rot} (d)	0.54095 ± 0.00003	4.578 ± 0.006
$v \sin i$ (km s^{-1})	72.6 ± 0.1	4.82 ± 0.16
v_{rad} (km s^{-1})	17.5 ± 0.1	8.34 ± 0.1
i ($^{\circ}$)	46 ± 10	42 ± 10
Distance (pc)	63.9 ± 1.4^a	46.27 ± 0.19^a
Ω_{eq} (rad d^{-1})	11.6199 ± 0.0005	–
$d\Omega$ (rad d^{-1})	0.0098 ± 0.0014	–

^aGaia Collaboration et al. (2016, 2018).

values derived in Section 3.1). For TWA 8A, we fixed the $v \sin i$ to 4.82 km s^{-1} , as this was determined by direct spectral fitting in Section 5 and is more accurate than that derived from ZDI.

4.1 Brightness and magnetic imaging

The observed LSD profiles for TWA 6 and TWA 8A, as well as our fits to data, are shown in Fig. 5. For TWA 6, we obtain a reduced chi-squared χ_r^2 fit equal to 1 (where the number of fitted data points is equal to 4312, with simultaneous fitting of both Stokes I and V line profiles). For TWA 8A, the low $v \sin i$ means that there is little modulation of the Stokes I line profiles, with the strong magnetic fields causing significant Zeeman broadening of the lines. Indeed, we are able to model the Stokes I line profiles sufficiently well using a stellar model with a homogeneous surface brightness, with our fits to the Stokes V line profiles yielding $\chi_r^2 = 1.04$ (for 930 fitted data points). We note that, given the substantially larger $v \sin i$ of TWA 6 as compared to TWA 8A, combined with more complete phase coverage, the reconstructed maps of TWA 6 have an effective resolution around 10 times higher.

The brightness map of TWA 6 includes both cool spots and warm plages (see Fig. 6), with no true polar spot, but rather a large spotted region centred around 60° latitude (centred around phase 0.6), with the majority of plages at a similar latitude on the opposing hemisphere. These features introduce significant distortions to the Stokes I profiles (see Fig. 5), introducing large RV variations (with maximum amplitude 6.0 km s^{-1} ; see Section 6). Overall, we find a spot and plage coverage of $\simeq 17$ per cent (10 and 7 per cent for spots and plages, respectively), similar to that found for V819 Tau, V830 Tau (Donati et al. 2015), and Par 2244 (Hill et al. 2017).

Note that the estimates of spot and plage coverage should be considered as lower limits only, as Doppler imaging is mostly insensitive to small-scale structures that are evenly distributed over the stellar surface (hence the larger minimal spot coverage assumed in Section 3.1 to derive the location of the stars in the H–R diagram).

4.2 Magnetic field imaging

Using our imaging code, we have reconstructed the magnetic fields of our target stars using both poloidal and toroidal fields, each

expressed using an SH expansion, with ℓ and m denoting the mode and order of the SH (Donati et al. 2006). For a given set of complex coefficients $\alpha_{\ell,m}$, $\beta_{\ell,m}$, and $\gamma_{\ell,m}$ (where $\alpha_{\ell,m}$ characterizes the radial field component, $\beta_{\ell,m}$ the azimuthal and meridional components of the poloidal field term, and $\gamma_{\ell,m}$ the azimuthal and meridional components of the toroidal field term), one can construct an associated magnetic image at the surface of the star, and thus derive the corresponding Stokes V data set. Here, we carry out the inverse, where we reconstruct the set of coefficients that fit the observed data.

For TWA 6, our reconstructed fields presented in Fig. 7 are limited to SH expansions with terms $\ell \leq 20$. Given the high $v \sin i$ of TWA 6 (combined with good phase coverage), we are able to resolve smaller scale magnetic fields, and indeed such a large number of modes are required to fit the observed Stokes V signatures. We note, however, that including higher order terms (>20) only marginally improves our fit. Such high-degree modes indicate that the magnetic fields in TWA 6 concentrate on smaller, more compact spatial scales. In contrast, our fits to the Stokes V observations of TWA 8A only require terms up to $\ell \leq 10$, with higher order terms providing only a marginal improvement. Hence, the magnetic field of TWA 8A is concentrated at larger spatial scales.

The reconstructed magnetic field for TWA 6 is split almost evenly between poloidal and toroidal components (53 and 47 per cent, respectively), with a total magnetic energy $\langle B \rangle = 840 \text{ G}$, where $\langle B \rangle$ is given by

$$\langle B \rangle = \iint_{\theta, \phi} (B_{\alpha}^2 + B_{\beta}^2 + B_{\gamma}^2)^{1/2} d\theta d\phi. \quad (3)$$

The poloidal field is mostly axisymmetric (49 per cent), with the largest fraction of energy (58 per cent) in modes with $\ell > 3$, and with 30 per cent of energy in the dipole mode ($\ell = 1$, with a field strength of 550 G). On large scales, the poloidal component is tilted at 35° from the rotation axis (towards phase 0.34). The toroidal component is also mostly axisymmetric, with the largest fraction of energy (68 per cent) in modes with $\ell > 3$, and with 17 per cent of energy in the octupole ($\ell = 3$) mode. These components combine to generate an intense field of $\geq 2 \text{ kG}$ at 45° latitude around phase 0.50–0.75 and 0.20–0.35, as well as an off-pole 2 kG spot at phase 0.75. We note that the large-spotted region reconstructed in

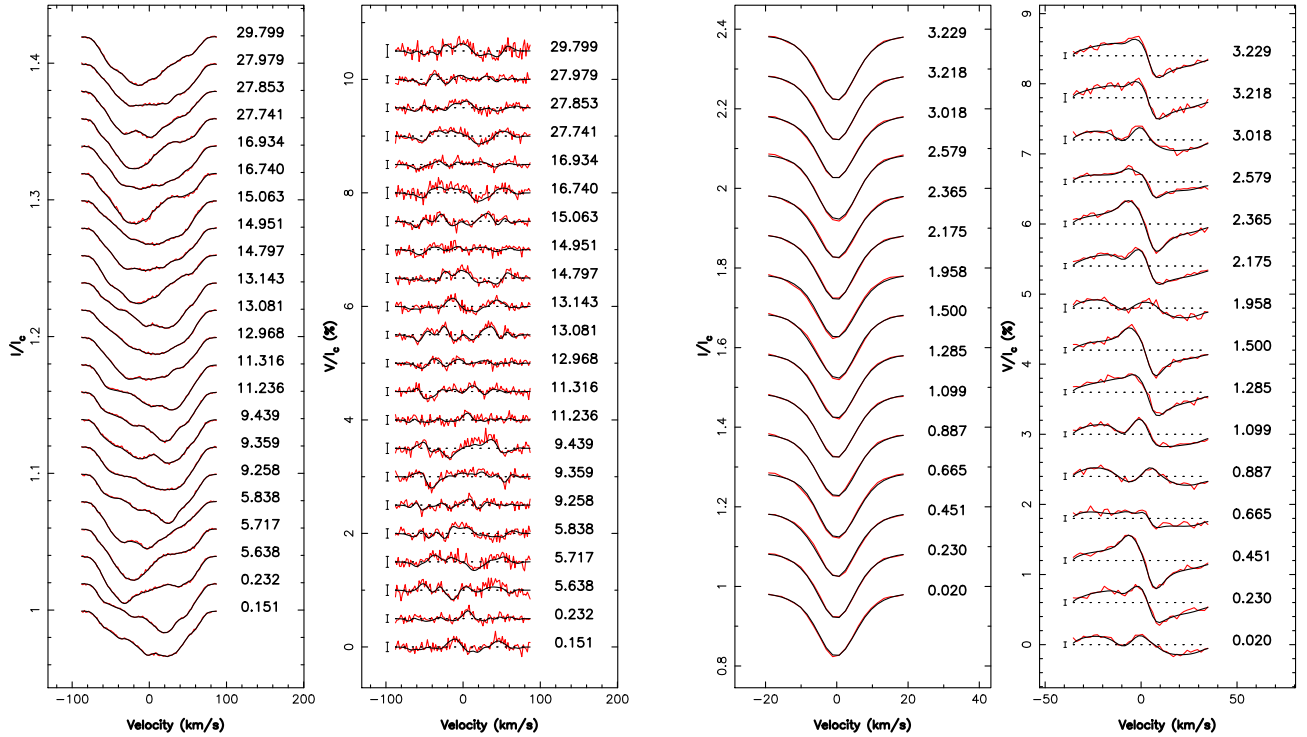


Figure 5. Maximum-entropy fit (thin red line) to the observed (thick black line) Stokes I (first and third panels) and Stokes V (second and fourth panels) LSD photospheric profiles of TWA 6 (first two panels) and TWA 8A (last two panels). Note that for TWA 8A the velocity scales are different. Rotational cycles are shown next to each profile. This figure is best viewed in colour.

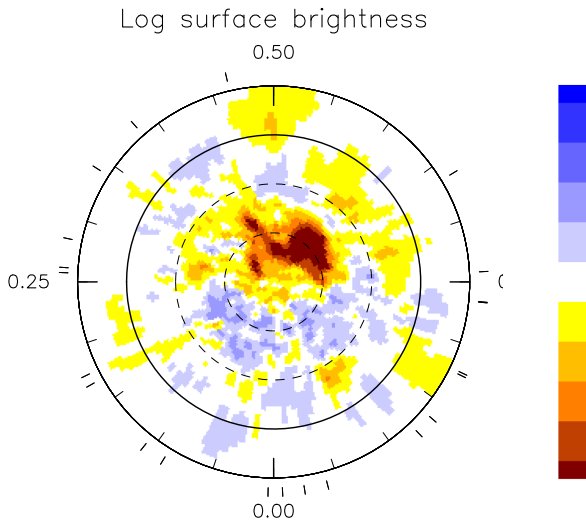


Figure 6. Map of the logarithmic brightness (relative to the quiet photosphere) at the surface of TWA 6. The star is shown in flattened polar projection down to latitudes of -30° , with the equator depicted as a bold circle, and 30° and 60° parallels as dashed circles. Radial ticks indicate the phases of observation. This figure is best viewed in colour.

the brightness map (around 60° latitude at phase 0.6, see Fig. 6) aligns well with these intense fields, suggesting that they are related.

In the case of TWA 8A, the reconstructed field is 71 per cent poloidal and 29 per cent toroidal, with a total unsigned flux of 1.4 kG, and with a magnetic filling factor of $f_v = 0.2$ (where f_v is equal to the fraction of the stellar surface that is covered by the

mapped magnetic field using Stokes V data). The poloidal field is mostly axisymmetric (70 per cent), with 16 per cent of the energy in the dipole ($\ell = 1$, with a field strength of 0.72 kG), 21 per cent in the quadrupole ($\ell = 2$), 18 per cent in the octupole ($\ell = 3$), and with the remaining 44 per cent of energy in modes with $\ell > 3$. On large scales (several radii from the star), the poloidal component may be approximated by an $B = 0.69$ kG aligned-dipole tilted at 20° from the rotation axis (towards phase 0.37). The toroidal component is mostly non-axisymmetric, with the majority of energy (55 per cent) in modes with $\ell > 3$, and with 21, 6 and 18 per cent in modes with $\ell = 3, 2$ and 1. These components combine to generate intense fields in excess of 2 kG in around phases 0.08, 0.42, and 0.75 on the stellar surface, centred around 20° latitude in the radial field component and around 35° in the meridional field component. Given the filling factor of $f_v = 0.2$, this suggests that surface magnetic fields can locally reach over 10 kG. Moreover, the high fraction of energy in high-order modes suggests that there are a large number of small-scale magnetic features, a conclusion also supported by the direct spectral fitting in Section 5.5.

In Fig. 8, we use the a potential field approximation (e.g. Jardine, Collier Cameron & Donati 2002) to extrapolate the large-scale field topologies of TWA 6 and TWA 8A. These topologies are derived solely from the reconstructed radial field components, and represent the lowest possible states of magnetic energy, providing a reliable description of the magnetic field well within the Alfvén radius (Jardine et al. 2013).

4.3 Surface differential rotation

The level of surface differential rotation of TWA 6 was determined in a similar manner as that carried out for other wTTSs (e.g. Skelly

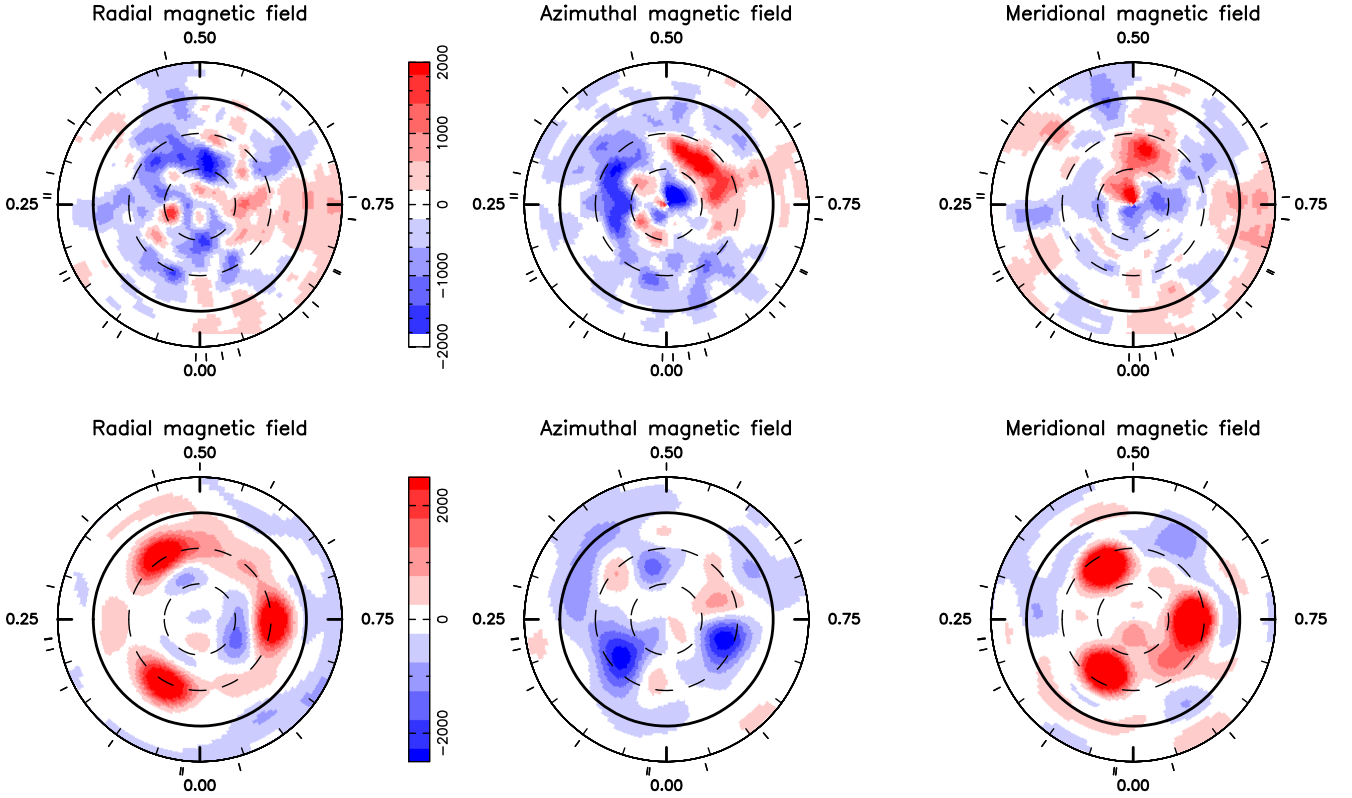


Figure 7. Map of the radial (left), azimuthal (middle), and meridional (right) components of the magnetic field \mathbf{B} at the surface of TWA 6 (top) and TWA 8A (bottom). Magnetic fluxes in the colourbar are expressed in G. Note that the magnetic filling factor for TWA 8A is $f_v = 0.2$. The star is shown in flattened polar projection as in Fig. 6. This figure is best viewed in colour.

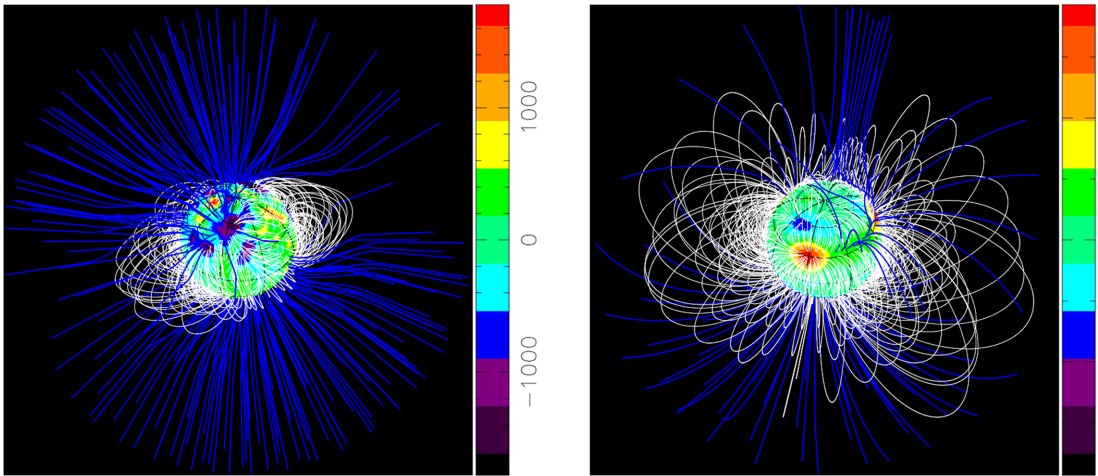


Figure 8. Potential field extrapolations of the radial magnetic field reconstructed for TWA 6 (left) and TWA 8A (right), viewed at phases 0.95 and 0.70, with inclinations of 45.6° and 31° , respectively. Open and closed field lines are shown in blue and white, respectively, whereas colours at the stellar surface depict the local values of the radial field (as shown in the left-hand panels of Fig. 7). The source surfaces at which the field becomes radial are set at distances of $2.6 R_*$ for TWA 6 and $10.7 R_*$ for TWA 8A, as these are close to the co-rotation radii (where the Keplerian orbital period equals the stellar rotation period, and beyond which the field lines tend to open under the effect of centrifugal forces, Jardine 2004), and are smaller than or similar to the Alfvén radii of $>10 R_*$ (Réville et al. 2016). This figure is best viewed in colour. Full animations may be found for both TWA 6 and TWA 8A at <https://imgur.com/hSkhYLT> and <https://imgur.com/AdKptUx>.

et al. 2008, 2010; Donati et al. 2014, 2015). Assuming that the rotation rate at the surface of the star varies with latitude θ as $\Omega_{\text{eq}} - d\Omega \sin^2 \theta$ (where Ω_{eq} is the rotation rate at the equator and $d\Omega$ is the difference in rotation rate between the equator and

the pole), we reconstruct brightness and magnetic maps at a fixed information content for many pairs of Ω_{eq} and $d\Omega$ and determine the corresponding reduced chi-squared χ_r^2 of our fit to the observations. The resulting χ_r^2 surface usually has a well-defined minimum to

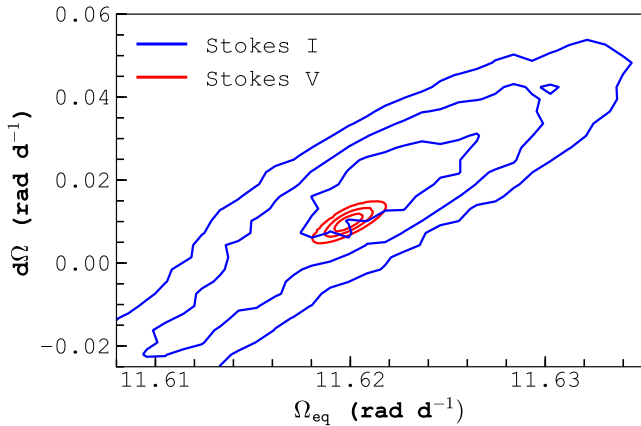


Figure 9. Variations of χ_r^2 as a function of Ω_{eq} and $d\Omega$ for TWA 6, derived from modelling of our Stokes I (red) and Stokes V (blue) LSD profiles at a constant information content. For both Stokes I and V , a clear and well defined parabola is observed, shown by the 1, 2, and 3σ ellipses (depicting 68.3, 95.5, and 99.7 per cent confidence levels, respectively), with the 3σ contour tracing the 5.5 per cent increase in χ_r^2 (or equivalently a χ^2 increase of 11.8 for 2156 fitted data points). This figure is best viewed in colour.

which we fit a parabola, allowing an estimate of both Ω_{eq} and $d\Omega$ (and their corresponding error bars).

Fig. 9 shows the χ_r^2 surface we obtain (as a function of Ω_{eq} and $d\Omega$) for both Stokes I and V for TWA 6. We find a clear minimum at $\Omega_{\text{eq}} = 11.6199 \pm 0.0005 \text{ rad d}^{-1}$ and $d\Omega = 0.0098 \pm 0.0014 \text{ rad d}^{-1}$ for Stokes I data (corresponding to rotation periods of $0.54073 \pm 0.00002 \text{ d}$ at the equator and $0.54118 \pm 0.00002 \text{ d}$ at the poles; see the left-hand panel of Fig. 9), with the fits to the Stokes V data of $\Omega_{\text{eq}} = 11.622 \pm 0.004 \text{ rad d}^{-1}$ and $d\Omega = 0.018 \pm 0.011 \text{ rad d}^{-1}$ showing consistent estimates, though with larger error bars (the right-hand panel of Fig. 9). We note that both these periods are in excellent agreement with those found previously by Skelly et al. (2008) and Kiraga (2012).

For TWA 8A, we were able to constrain the rotational period to $4.578 \pm 0.006 \text{ d}$ (corresponding to $\Omega_{\text{eq}} = 1.3724 \pm 0.0019 \text{ rad d}^{-1}$), in good agreement with the photometric period of 4.638 d found by Kiraga (2012). However, given that the observations span only ~ 3 rotation cycles, the recurrence of profile distortions across different latitudes is severely limited, and so we were unable to constrain surface shear. Hence, for our fits with ZDI, we have assumed solid body rotation.

5 MAGNETIC FIELD STRENGTH FROM INDIVIDUAL LINES

TWA 8A has a very strong photospheric magnetic field that can be detected in some individual lines, allowing direct spectral fitting to derive the strength of the magnetic field. As this is not the case for TWA 6, it is not included in the following analysis. For TWA 8A, Stokes V signatures are visible in over 20 lines, mostly redwards of 8000 \AA where the S/N is largest. Of particular interest are a set of 11 strong Ti I lines between 9674 and 9834 \AA ; 10 of which are detected in Stokes V and 1 of which has a Landé factor of 0 (9743.6 \AA , see Fig. 10). These atomic lines have minimal blending from molecular lines, and while there is a some blending from telluric lines, it can be corrected. These lines have the added advantage that all but two of them are from the same multiplet, which mitigates the impact of some systematic errors (e.g. errors in T_{eff}) on our measurements of

the magnetic field. A detailed description of these lines is given in Table C1.

5.1 Telluric correction

Before a detailed analysis of the Stokes I spectra may be carried out, we must first correct for the large number of telluric water lines present between 9670 and 9840 \AA . Telluric lines are not expected to produce circular polarisation, and we see no indication of them in Stokes V , hence we conclude that their impact on the Stokes V spectrum is negligible.

As we did not expect to detect magnetic fields in individual telluric blended lines, we did not observe a hot star for telluric calibration. Fortunately, on some nights, other programmes with ESPaDOnS at the CFHT observed the hot stars HD 63401 (PI: J.D. Landstreet) and HD 121743 (PI: G.A. Wade). HD 63401 is a 13500 K , Bp star (e.g. Bailey 2014) and HD 121743 is a 21000 K , B star (e.g. Alecian et al. 2014), with both stars having virtually no photospheric lines in the wavelength range of interest, apart from Paschen lines. Our observations of TWA 8A on the nights of March 25 to April 1, as well as April 5 and 6, had suitable telluric reference observations that were sufficiently close in time and obtained under sufficiently similar conditions.

The telluric reference spectra were first continuum normalized by fitting low-order polynomials through carefully selected continuum regions, then dividing by those polynomials, independently for each spectral order. The telluric reference spectra were then scaled in the form I^a , where I is the continuum normalized spectrum and a the scaling factor. The scaling factor a and the radial velocity shift for the telluric lines were determined by fitting the modified reference spectrum to telluric lines of the science spectrum through χ^2 minimization. Telluric lines around the photospheric lines of interest (~ 9650 – 9850 \AA) were included, as well as some telluric lines in the range of 9300 – 9500 \AA where there are fewer strong photospheric lines. The science spectrum was then divided by the scaled shifted telluric spectrum. An example spectrum before and after telluric correction is shown in Fig. 10.

5.2 Spectrum synthesis

To constrain the strength of the photospheric magnetic field, we have modelled individual lines in the Stokes I and V spectra of TWA 8A. Furthermore, as one of the Ti I lines has a Landé factor of 0 and is narrower in Stokes I as compared to the other Ti I lines, the magnetic field can also be strongly constrained by the Stokes I spectrum.

To generate synthetic spectra, we used the ZEEMAN spectrum synthesis program (Landstreet 1998; Wade et al. 2001; Folsom et al. 2012). This program includes the Zeeman effect and performs polarized radiative transfer in Stokes $IQUV$. The code uses plane-parallel model atmospheres and assumes LTE, and produces disc-integrated spectra. ZEEMAN includes quadratic Stark, radiative, and van der Waals broadening, as well as optional microturbulence (v_{mic}) and radial-tangential macroturbulence. A limitation of the code for use in very cool stars is that it does not include molecular lines, or calculations of molecular reactions in the abundances for atomic species. The Ti I lines in the 9674 – 9834 \AA region are blended with a few very weak molecular lines, and so ZEEMAN can produce accurate spectra for this region; however, most of the spectral region bluewards of this is problematic.

For input to the code, we used MARCS model atmospheres (Gustafsson et al. 2008) and atomic data taken from VALD

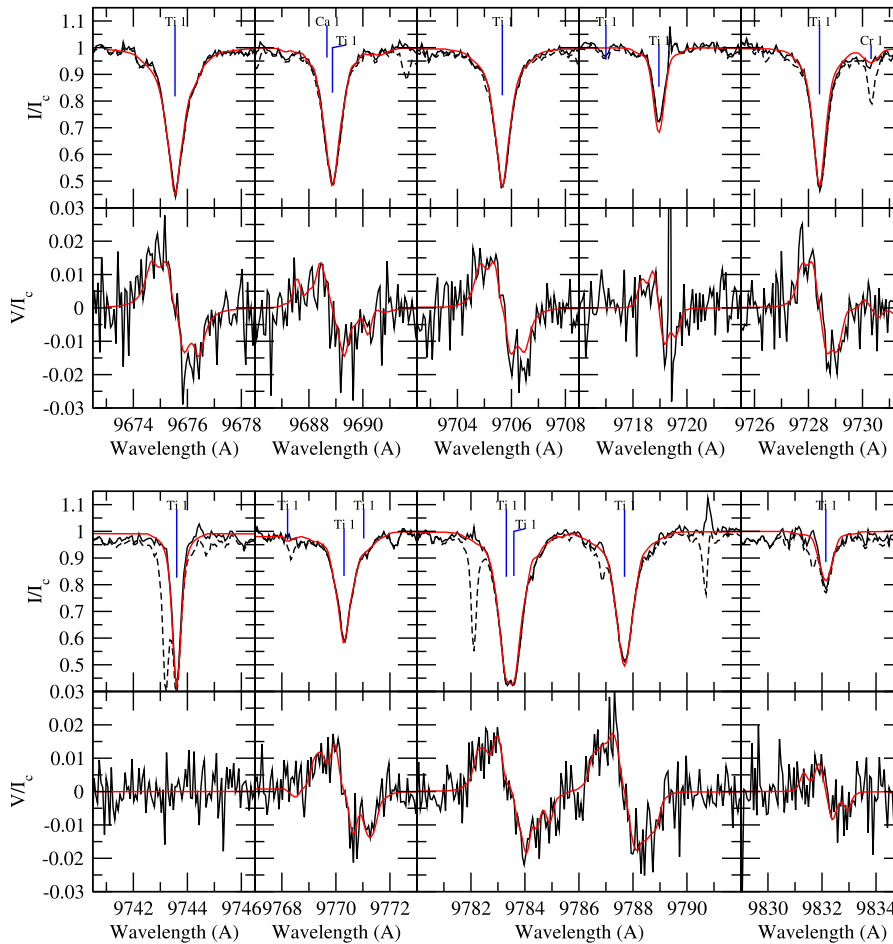


Figure 10. Detections of Zeeman broadening in the observation of TWA 8A on March 27. The panels show the Stokes I spectrum at the top and the corresponding Stokes V spectrum below for the full set of lines used in our fits (see Table C1). Dashed lines show the observation before telluric correction, and solid lines show the spectrum after telluric correction. Overplotted in a red solid line is our best fit using our third model to fit both Stokes I and V simultaneously.

(Ryabchikova et al. 2015; see Table C1 for the properties of the atomic lines). VALD data for these particular Ti I lines were also used by Kochukhov & Lavail (2017) for a similar analysis and were deemed reliable. Additionally, we can reproduce these Ti lines with near-solar abundances, implying that the oscillator strengths are likely close to correct.

To model the magnetic field of TWA 8A, we adopted a uniform radial magnetic field. While this is an unrealistically simple magnetic geometry, the ZDI analysis found the magnetic geometry to be more complex than a simple dipole. Therefore, we leave the geometric analysis to ZDI and adopt the simplest possible geometry here to avoid additional weakly constrained geometric parameters. Furthermore, since this analysis is applied to individual observations, a full magnetic geometry cannot be reliably derived. The model we implement here includes a combination of magnetic field strengths B , each with their own filling factor f , with the sum of the filling factors (including a region of zero field) equal to unity.

We fit synthetic spectra using a Levenberg–Marquardt χ^2 minimization routine (similar to Folsom et al. 2012, 2016), with the radial magnetic field strengths and filling factors as optional additional free parameters. The code was updated to allow fitting observed Stokes I spectra, V spectra, or I and V simultaneously, with wavelength ranges carefully set around the lines of interest. In order to place

uncertainties on the fitting parameters, we use the square root of the diagonal of the covariance matrix, as is commonly done. This is then scaled by the square root of the reduced χ^2 to very approximately account for systematic errors. These formal uncertainties may still be underestimates, and a further consideration of uncertainties is discussed in Section 5.5.

5.3 Fitting the Stokes I spectrum

Our initial fits were carried out with the observation on March 27 since the Stokes V LSD profile for this night has one of the simplest shapes, indicating a more uniform magnetic field in the visible hemisphere.

Measurements of magnetic fields in Stokes I spectra are constrained by both the width and the desaturation of lines with different Landé factors. Fitting the Stokes I spectrum to determine magnetic field strengths requires constraints on several other stellar parameters that influence line width and depth. Here, we adopt the T_{eff} and $\log g$ values derived in Section 3 (see Table 2). Since our choice of lines is dominated by one multiplet, adopting these values is a small source of uncertainty. We note that these lines are not well adapted to constraining T_{eff} and $\log g$ spectroscopically. We include $v \sin i$ and v_{mic} as free parameters in the fit, since they can

play an important role in line shape and strength, and can only be determined spectroscopically. v_{mic} is constrained by desaturation of strong (on the curve of growth) lines and, given the lack of weak lines in our spectral range, is determined with only a modest accuracy by different degrees of desaturation of different strong lines. Macroturbulence is assumed to be zero, since it is likely much smaller than the $v \sin i$ of $\sim 5 \text{ km s}^{-1}$. Ti abundance is included as a free parameter, however, we caution the reader that this may not provide reliable results, as the code neglects the fraction of Ti bound in molecules. Nevertheless, this free parameter is necessary to avoid the code fitting line strength entirely by varying magnetic field and v_{mic} .

When fitting the spectra of TWA 8A we adopted three main models, each of increasing complexity, to better constrain the nature of the magnetic field. These three models (described below) are used to fit Stokes I spectra only, Stokes V only, and both Stokes I and V simultaneously.

Our first model consists of fitting the Stokes I spectrum using just one magnetic region with a corresponding filling factor, yielding a best-fitting magnetic field strength of $B_1 = 5.65 \pm 0.10 \text{ kG}$ with a filling factor $f_1 = 0.597 \pm 0.016$, but at a reduced χ^2 of 19.6. Fits with f fixed to 1 consistently fail to reproduce the line shape, with a core that is far too wide and with wings that are too narrow, implying that only a fraction of the star is covered by very strong magnetic fields.

Our second model increases the number of free parameters by including two magnetic regions and filling factors, achieving a visibly much better fit with a reduced χ^2 of 12.9, and with field strengths of $B_1 = 4.71 \pm 0.08 \text{ kG}$ with $f_1 = 0.648 \pm 0.015$, and $B_2 = 15.61 \pm 0.25 \text{ kG}$ with $f_2 = 0.133 \pm 0.007$. This second model does a better job of simultaneously reproducing the narrow core and broad wings of the magnetically sensitive lines, although the high-field strength region produces a sharper change in the shape of the wings than seen in the observation, implying that the star has a more continuous distribution of magnetic field strengths than our model.

Our third model again increases the number of free parameters to improve the fit. However, rather than add additional sets of magnetic field strengths and filling factors, which may become more poorly conditioned or not converge well, we instead adopt a grid of fixed magnetic field strengths with filling factors as free parameters (in a similar way to e.g. Johns-Krull, Valenti & Koresko 1999; Johns-Krull, Valenti & Saar 2004). This provides an approximate distribution of magnetic field strengths on the visible hemisphere of the star. Using our third model for fitting Stokes I only, we use bins of 0, 2, 5, 10, 15, and 20 kG. Bins of $\sim 5 \text{ kG}$ allow for smooth model line profiles, and so smaller bins (that would be less well constrained) are not necessary. Adding bins above 20 kG improves the χ^2 fit by a small but formally significant amount. However, the impact on the synthetic line is small and only affects the far wings of the line in Stokes I . Small changes in the far wings of the line are most vulnerable to systematic errors, such as weak lines that are not accounted for, errors in the telluric correction, errors in continuum normalization, or very weak fringing, all of which could approach the strength of the line this far into the wing. Thus, we limit the magnetic field to 20 kG and caution that even for this bin the filling factor may be overestimated. The resulting best-fitting parameters for Stokes I only for March 27 using our third model is presented in Table 3, with a reduced χ^2 of 10.6.

Yang, Johns-Krull & Valenti (2008) studied TWA 8A and derived some magnetic quantities based on Stokes I observations in the IR. They adopted literature values for the stellar parameters of $T_{\text{eff}} =$

3400 K, $\log g = 4.0$ and $v \sin i = 4.0 \text{ km s}^{-1}$. Their ‘Model 1’ corresponds to our first model with one filling factor and magnetic field strength. They report only the product of their filling factor and magnetic field strength as 2.3 kG, which is close to our value for March 27 of $3.37 \pm 0.11 \text{ kG}$, although not within uncertainty. Their ‘Model 2’ corresponds to our second model with two filling factors and magnetic field strengths. They report the quantity $\langle |B_f| \rangle = 2.7 \text{ kG}$, which is comparable but again not consistent with our value of $5.13 \pm 0.14 \text{ kG}$. The ‘Model 3’ of Yang et al. (2008) is closest to our third model with a grid of filling factors, although they only fit filling factors for field strengths of 2, 4, and 6 kG. They report $\langle |B_f| \rangle = \sum_i B_f i$ of 3.3 kG. The equivalent value from our fit is $\langle |B_f| \rangle = 5.90 \pm 0.44 \text{ kG}$, which is again inconsistent. We note that, if we perform our fit using the three bins of 2, 4, and 6 kG used by Yang et al. (2008), we find $\langle |B_f| \rangle = 3.96 \pm 0.38 \text{ kG}$. While this is much closer to their ‘Model 3’ results, we find that the fit to our data is much worse in the wings of the lines, so we consider this model to be less accurate for our spectra. The IR spectra of Yang et al. (2008) had a much lower S/N than our observations, and so the wings of the lines may not have been detected as clearly as in our spectra. Indeed, the very strong magnetic field with a very small filling factor necessary to fit the wings of our magnetically sensitive lines is likely the cause of the difference between our results, as well as intrinsic variability of the field.

5.4 Fitting the Stokes V spectrum

In order to fit the Stokes V spectrum we adopt the best fit $v \sin i$, v_{mic} and Ti abundance from fitting Stokes I with our third model, since these parameters cannot be well constrained from V spectra (see Table 3).

When directly fitting the Stokes V spectrum, it becomes immediately apparent that a filling factor (much less than unity) is necessary. To produce Stokes V profiles with the widths of the observed lines, a very strong magnetic field is necessary. However, to reproduce the amplitudes of the Stokes V profiles, a weaker field is necessary, or a very strong field covering a small portion of the star. This can be easily seen by comparing the widths of the observed Stokes I and V profiles (see Fig. 5) and noting that the V profiles remain stronger in the far wings compared to the I profiles.

Fitting the Stokes V profiles with our first model yields a best fit of $B_1 = 7.09 \pm 0.19 \text{ kG}$ and $f_1 = 0.081 \pm 0.004$, with a reduced χ^2 of 2.27. However, this provides a poor fit to the line profiles, in particular the outer and inner parts of the line cannot be well fit simultaneously. We find a much better fit when using our second model, with a reduced χ^2 of 1.58, and field strengths and filling factors of $B_1 = 4.70 \pm 0.19 \text{ kG}$ with $f_1 = 0.078 \pm 0.004$, and $B_2 = 14.94 \pm 0.26 \text{ kG}$ with $f_2 = 0.051 \pm 0.003$, implying $\langle |B_f| \rangle = 1.13 \pm 0.05 \text{ kG}$. The filling factors and $\langle |B_f| \rangle$ derived here are much smaller than those derived from Stokes I . Stokes V is sensitive to the sign of the line-of-sight component of \mathbf{B} , while Stokes I is sensitive to the magnitude of \mathbf{B} . The difference in filling factors is likely due to cancellation in V of nearby regions with opposite sign.

We also fit the Stokes V spectra with our third model, where our use of positive fields is still appropriate as the disc integrated field is positive for March 27, and indeed at all other phases. Our fit yields a reduced χ^2 of 1.56, where the parameters are summarized in Table 3. The improvement in the fit using our third model is modest compared to the first and second models, but it is clearly better visually, with a formally significant improvement of nearly 3σ . We note that the distribution of filling factors is quite different

Table 3. Best-fitting parameters from direct spectral fitting of TWA 8A. The first and second columns, respectively, give the results of fitting Stokes I and V separately (using our third model) for the spectrum taken on 2015 March 27. Parameters with no error bars for the V fit were held fixed. The third column shows the results of fitting Stokes I and V simultaneously, where we present the mean over the 10 nights that could be reliably telluric corrected, with error bars given as the standard deviations. Values for fits to individual nights are presented in Table C2.

	Stokes I only 2015 March 27	Stokes V only 2015 March 27	Stokes I and V Mean
$v \sin i$ (km s $^{-1}$)	4.77 ± 0.23	4.77	4.82 ± 0.16
v_{mic} (km s $^{-1}$)	1.15 ± 0.08	1.15	1.08 ± 0.05
[Ti/H]	-7.006 ± 0.017	-7.01	-6.976 ± 0.022
0 G	0.002 ± 0.084	0.859 ± 0.014	0.017 ± 0.016
+2 kG	0.312 ± 0.067	0.023 ± 0.010	0.161 ± 0.016
+5 kG	0.483 ± 0.044	0.057 ± 0.007	0.245 ± 0.019
+10 kG	0.090 ± 0.020	0.016 ± 0.005	0.055 ± 0.004
+15 kG	0.060 ± 0.013	0.041 ± 0.005	0.047 ± 0.003
+20 kG	0.053 ± 0.011	0.004 ± 0.004	0.035 ± 0.004
-2 kG	-	-	0.155 ± 0.010
-5 kG	-	-	0.209 ± 0.007
-10 kG	-	-	0.041 ± 0.005
-15 kG	-	-	0.010 ± 0.003
-20 kG	-	-	0.027 ± 0.004
$\langle B_f \rangle$ kG	5.9 ± 0.4	1.2 ± 0.1	5.9 ± 0.2

from that of the Stokes I fit, with most of the surface having no magnetic field detected in Stokes V , and the remaining field lying more in the 5 and 15 kG bins.

Using our fits to approximate the longitudinal magnetic field (B_l), we have taken the line-of-sight component of the model magnetic field, averaged over the stellar disc and weighted by the brightness of the continuum, i.e.

$$B_{l,\text{syn}} = \sum_i \frac{\int I_c f_i B_i \cos(\theta) d\Omega}{\int I_c d\Omega}, \quad (4)$$

where f_i is the filling factor for component i , B_i is the purely radial magnetic field for that component, θ is the angle between the line of sight and the radial field. I_c is the continuum brightness at for a point on the disc (accounting for limb darkening), and the integral of $d\Omega$ is over the visible disk.

From equation (4), we derive $B_{l,\text{syn}} = 0.78 \pm 0.04$ kG and 0.82 ± 0.08 kG for our second and third models, respectively. These values agree to within their uncertainties, and are comparable to (but roughly 1.7 times larger than) the actual observed B_l values for this phase, as calculated from the LSD profiles (see Fig. 4). Indeed, if we calculate an observed B_l from just the Ti I 9705.66 Å line (using the telluric corrected spectrum), rather than an LSD profile, we find 0.88 ± 0.13 kG for March 27. Moreover, the behaviour of this Ti I line with rotational phase is consistent with the LSD profile, except that it shows a higher field strength. This implies that the signal in the Stokes V LSD profiles may not be adding perfectly coherently, producing a lower amplitude V profile. This is not surprising as, due to the very large field strength, Zeeman splitting patterns of individual lines begin to matter for the line profile shapes. Thus, simply scaling amplitudes by effective Landé factors is a less effective approximation for such strong fields.

5.5 Simultaneous fitting of Stokes I and V

As we detect magnetic fields in both Stokes I and V observations, our model should be able to reproduce these signatures simultaneously. This requires us to allow a combination of positive and negative magnetic fields, resulting in a cancellation of much of the signal in Stokes V while allowing for a large unsigned magnetic flux in

Stokes I . This is evident from the much smaller filling factor in our fits of Stokes V compared to our fits to Stokes I .

First, we performed simultaneous fits to Stokes I and V using a simple model with three magnetic regions – two with positive fields and one with a negative field. A model with one positive field and one negative field is insufficient to reproduce the shapes of the Stokes I or V line profiles. For this simple model, the best-fitting magnetic parameters are $B_1 = +4.76 \pm 0.07$ kG with $f_1 = 0.360 \pm 0.007$, $B_2 = -5.05 \pm 0.09$ kG with $f_2 = 0.282 \pm 0.007$, and $B_3 = +15.92 \pm 0.20$ kG with $f_3 = 0.098 \pm 0.004$ (with $v \sin i = 5.26 \pm 0.17$ km s $^{-1}$, $v_{\text{mic}} = 1.00 \pm 0.06$ km s $^{-1}$ and [Ti/H] = -6.947 ± 0.013). This fit gives a reduced χ^2 of 7.94, and fits the I spectrum similarly well to our best model from fitting Stokes I only (see above), although it is too strong in the wings of V , implying that there should be additional cancellation. This model implies a total $\langle |B_f| \rangle$ of 4.70 kG, and a synthetic $B_{l,\text{syn}}$ (allowing for cancellation) of 1.28 kG, although (as noted) this is likely too large.

Using our third model (with a grid of magnetic field strengths and filling factors, see above), we again require both negative and positive magnetic fields. As with fitting only Stokes I or V , we use bins of 0 G, ± 2 kG, ± 5 kG, ± 10 kG, ± 15 kG, and ± 20 kG, for a total of 11 bins. The results of our fit with this model, with 11 filling factors as well as $v \sin i$, v_{mic} , and [Ti/H], are presented in Table C2, with a reduced χ^2 of 6.33 – clearly an improvement over the simple three magnetic-region model. Our fit to the observation taken on March 27 is shown in Fig. 10, showing a good fit to both Stokes I and V spectra, including matching the width of the magnetically insensitive line with a Landé factor of 0.

A summation of the filling factors for bins with the same $|B|$ yields a very similar distribution to that for the fit to Stokes I only, with differences much smaller than the formal uncertainties. This can be understood as Stokes I is sensitive to the total magnetic field strength but not the orientation of the magnetic field. Similarly, the difference between filling factors for bins with the same $|B|$ but opposite sign produces a distribution very similar to that of the fit to Stokes V only. This can also be understood since Stokes V is sensitive to the line-of-sight component of the magnetic field only, with the spatially unresolved (within the same model pixel) components of opposite orientation cancelling out. For our observation on

March 27, we find a total $\langle |B_f| \rangle = 5.71 \pm 0.22$ kG and $B_{l,\text{syn}} = 0.78 \pm 0.15$ kG. This $\langle |B_f| \rangle$ is consistent with our fit of only Stokes I with our third model, and $B_{l,\text{syn}}$ is consistent with our fit of only Stokes V .

Over the rotation of TWA 8A, this set of results shows $B_{l,\text{syn}}$ to range from 640 ± 150 to 840 ± 140 G, with $\langle |B_f| \rangle$ ranging from 5.71 ± 0.22 to 6.36 ± 0.22 kG, and varying coherently with rotation phase.

Given the high S/N of our observations, the results we present here may be limited by systematic errors, and our uncertainties may be underestimated. To investigate the impact of uncertainties in T_{eff} and $\log g$, we re-fit the observation on March 27 with these two parameters changed by $\pm 1\sigma$. The change in T_{eff} produces at most a change of 0.5σ in the other parameters, and often smaller changes than that, and so we conclude that the uncertainty on T_{eff} has a minor contribution to the total uncertainty. Changing $\log g$ by 1σ has a large impact on $v \sin i$ and [Ti/H] ($4-5\sigma$) and on v_{mic} (2σ), although it has a much smaller impact on the magnetic filling factors of only $\sim 1\sigma$, rising to 3σ for the 2 kG and 5 kG bins when $\log g$ is decreased by 1σ . In that case, the filling factor shifts from the 2 kG bin into the 0 and 5 kG bins, underscoring the uncertainty of the 2 kG bin. The relatively large uncertainty in $\log g$ changes the line broadening, but does so independently of Landé factor, and so $v \sin i$ and v_{mic} are more sensitive to $\log g$ than filling factors. It is possible that our v_{mic} is an overestimate, since typical v_{mic} values for PMS M-dwarfs are not well known. To estimate an upper limit on this uncertainty, we re-ran the fit with $v_{\text{mic}} = 0$, finding that the best-fitting $v \sin i$ decreases by 1 km s^{-1} , that [Ti/H] increases by 0.1 dex, and that filling factors generally change by less than 1σ (except for the 10 kG bin which decreases by 2σ). From these tests, we conclude that our formal uncertainties may be underestimated by a factor of $\lesssim 2$, mostly due to the large uncertainty in $\log g$ and the (potentially) larger systematic errors on the filling factors for the 2 and 20 kG bins.

Having established an analysis method for the observation of March 27 using our third model to fit both Stokes I and V , we performed this analysis on all observations for which we could perform reliable telluric correction, providing us with 10 sets of results, shown in Table C2. Taking an average over all 10 observations, we find a mean magnetic field strength of $\langle |B_f| \rangle = 5.9 \pm 0.2$ kG, where the amount of magnetic energy in each bin is shown in Table 3. The standard deviation of these results is close to the mean uncertainty for all parameters, suggesting that our formal uncertainties account well for random errors, with the larger standard deviation likely due to the rotational modulation.

In Fig. 11, we compare the magnetic field strength distribution on TWA 8A as determined by our ZDI maps in Section 4.2, to our direct spectral fitting here. As our ZDI map has a continuous distribution of field strengths, we have created histograms using the same bins as that for the direct spectral fitting, allowing for a direct comparison of recovered field strengths. For Stokes I , we find that 75 per cent of the field strength recovered by ZDI is in the 2 kG bin, with a 15 per cent in the 5 kG bin, and 9 per cent at higher field strengths. In comparison, direct spectral fitting yields 32 per cent of the magnetic field to be 2 kG, with almost 46 per cent in the 5 kG bin, and with 22 per cent of fields in the 10, 15 and 20 kG bins. For Stokes V , the line profiles are sensitive to the sign of the line-of-sight component of \mathbf{B} , and so there is likely significant cancellation of fields of opposite polarity. Hence, our fits to Stokes V LSD profiles with ZDI recover only the uncanceled magnetic fields. Therefore, for comparison to direct spectra fitting, we must subtract the filling factors determined for the negative fields from the positive fields,

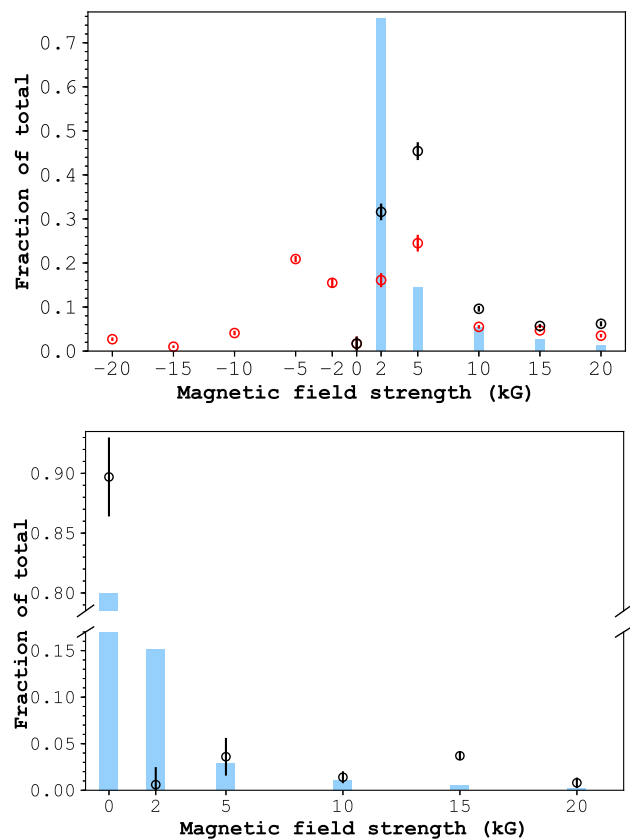


Figure 11. The distribution of surface magnetic field strengths for TWA 8A, as determined from ZDI and direct spectral fitting of Ti I lines. Blue bars show the fraction of the total mapped magnetic field strength from ZDI, for fields of a given bin. Top panel: Comparison between the magnetic field strengths determined from fitting Stokes I data. Red circles show the mean filling factors for each field strength using our third model to simultaneously fit Stokes I and V spectra (see Section 5.5). Black circles show the combined filling factors for both the positive and negative fields. Thus, one can directly compare the recovered field strengths for Stokes I data from ZDI and direct spectral fitting by comparing the blue bars and the black circles, respectively. One can see that a significantly larger fraction of higher strength fields are recovered by direct spectral fitting, as compared to that of ZDI (see the discussion in Section 5.5). Bottom panel: Comparison between the magnetic field strengths determined from fitting Stokes V data. Black circles show the resulting filling factors after subtracting the contributions of the negative fields from those of the positive fields. Note that the filling factor for the 0 kG bin is calculated as 1 minus the sum of the non-zero bins, so the sum total of all fractions is unity. As Stokes V profiles are sensitive to the sign of the line-of-sight component of \mathbf{B} , significant cancellation of fields may occur, and so we must compare our fits with ZDI to Stokes V profiles, to these black circles. In this case, we see that ZDI recovers a similar fraction of field strengths for the 5 and 10 kG bins, significantly less for the 15 and 20 kG bins, but significantly more for the 2 kG bin (see discussion in Section 5.5). ZDI also recovers a smaller fraction for the 0 kG bin, likely due to degeneracy between the 0 and 2 kG fields in direct spectral fitting.

yielding the fraction of uncanceled fields that could be fit with ZDI. For ZDI, we find that 80 per cent of the surface has a 0 G field, with 15 per cent of the field in the 2 kG bin, 3 per cent in the 5 kG bin, 1 per cent in the 10 kG bin, and with 1 per cent at higher field strengths. In comparison, for direct spectral fitting we find that less than 1 per cent of the field is 2 kG, with 3.6 per cent of the field at 5 kG, 1.4 per cent at 10 kG, and with 4.5 per cent at higher field strengths. Thus, with ZDI we recover most of the

magnetic flux up to 10 kG, but are not as sensitive to fields higher than this. Moreover, our results demonstrate that we underestimate the fraction of high field strengths using the ZDI technique with LSD profiles Stokes V spectra. As mentioned previously, this may be due to the signal in the Stokes VLSD profiles not adding perfectly coherently, as variations in line splitting patterns cause variations in line shapes, and so scaling amplitudes by effective Landé factors is less accurate. Moreover, there may be significant cancellation in Stokes V profiles, as it is sensitive to the sign of the line-of-sight component of \mathbf{B} . The recovery of small-scale, high-field-strength features would likely be improved if linear polarization spectra (Stokes Q and U) were included in the ZDI modelling, and would likely increase the recovered total magnetic field energy (see Rosén, Kochukhov & Wade 2015).

6 FILTERING THE ACTIVITY JITTER

As well as characterizing magnetic fields of wTTSs, the MaTYSSSE programme also aims to detect close-in giant planets (called hot Jupiters, hJs) to test planetary formation and migration mechanisms. In particular, characterizing the number and position of hJs will allow us to quantitatively assess the likelihood of the disc migration scenario, where giant planets form in the outer accretion disc and then migrate inwards until they reach the central magnetospheric gaps of cTTSs (see e.g. Lin, Bodenheimer & Richardson 1996; Romanova & Lovelace 2006). Given that we map the surface brightness of the host star, we are able to use our fits to the observed data to filter out the activity-related jitter from the RV curves (where the RV is measured as the first-order moment of the LSD profile; see Donati et al. 2014, 2015). After subtraction of the RV jitter, we may look for periodic signals in the RV residuals to reveal the presence of hJs. Indeed, this method has so far yielded two detections of hJs in the MaTYSSSE sample, around both V830 Tau (Donati et al. 2015, 2016, 2017) and TAP 26 (Yu et al. 2017).

For TWA 6, the unfiltered RVs have an rms dispersion of 3.8 km s^{-1} . The predicted RV due to stellar activity and the filtered RVs are shown in Fig. 12. We find that RV residuals exhibit an rms dispersion of $\sim 0.20 \text{ km s}^{-1}$, with a maximum amplitude of 0.51 km s^{-1} . This is well above the intrinsic RV precision of ESPaDOnS (around 0.03 km s^{-1} ; e.g. Moutou et al. 2007; Donati et al. 2008); however, given the high $v \sin i$, the accuracy of the filtering process is somewhat reduced, with an intrinsic uncertainty of around 0.1 km s^{-1} . Indeed, we find no significant peaks in a periodogram analysis, and so we find that TWA 6 is unlikely to host a hJ with an orbital period in the range of what we can detect (i.e. not too close to the stellar rotation period or its first harmonics; see Donati et al. 2014). We find a 3σ error bar on the semi-amplitude of the RV residuals equal to 0.19 km s^{-1} , translating into a planet mass of $\simeq 3.1 M_{\text{Jup}}$ orbiting at $\simeq 0.1 \text{ au}$ (assuming a circular orbit in the equatorial plane of the star; see Fig. 13).

For TWA 8A, the unfiltered RVs have an rms dispersion of 0.13 km s^{-1} . Given that the surface brightness of TWA 8A is compatible with that of a homogeneous star, we were unable to filter the RVs in the same manner. However, the measured RVs (shown in Fig. 12) do display a clear periodic signal that is equal to the stellar rotation period, implying that there are starspots on the surface, even though the modulation of the line profiles is minimal.

7 SUMMARY AND DISCUSSION

We report the results of our spectropolarimetric observations collected with ESPaDOnS at CFHT of two wTTSs, namely TWA 6 and

TWA 8A, in the framework of the international MaTYSSSE Large Program. Our spectral analysis reveals that the two stars have quite different atmospheric properties, with photospheric temperatures of 4425 ± 50 and $3690 \pm 130 \text{ K}$ and logarithmic gravities (in cgs units) of 4.5 ± 0.2 and 4.3 ± 0.3 . The stars are significantly different in mass, with TWA 6 being $1.0 \pm 0.10 M_{\odot}$ and TWA 8A being around half that at $0.45 \pm 0.10 M_{\odot}$. Likewise, the radii are also different with $1.0 \pm 0.2 R_{\odot}$ for TWA 6 and $0.8 \pm 0.2 R_{\odot}$ for TWA 8A, viewed at inclinations of $46^{\circ} \pm 10^{\circ}$ and $31^{\circ} \pm 10^{\circ}$. Using the Siess et al. (2000) evolutionary models (for direct comparison to other MaTYSSSE and MaPP results), we estimate their ages to be 21 ± 9 and $11 \pm 5 \text{ Myr}$, with TWA 6 being mostly radiative, and TWA 8A being fully convective. We note that these masses, ages, and internal structures depend strongly on the adopted temperatures.

With a rotation period of $0.54095 \pm 0.00003 \text{ d}$, TWA 6 is the most rapidly rotating wTTS yet mapped with ZDI, and one of the fastest rotators in TWA (see de la Reza & Pinzón 2004). By contrast, TWA 8A has a much slower period of $4.578 \pm 0.006 \text{ d}$, which is very similar to the median period of 4.7 d of the TWA 1–13 group (Lawson & Crause 2005), and also more similar to that of other wTTSs such as V819 Tau ($P_{\text{rot}} = 5.53113 \text{ d}$, Donati et al. 2015), as well as Par 1379 ($P_{\text{rot}} = 5.585 \text{ d}$, Hill et al. 2017).

We find that neither TWA 6 nor TWA 8A has an IR excess up to $23.675 \mu\text{m}$. Hence, both stars have likely dissipated their circumstellar accretion discs, with either no accretion taking place, or with accretion occurring at an undetectable level, given that standard accretion-rate metrics based on the EWs of H α , H β , and He I D₃ are strongly affected by chromospheric emission.

The H α , H β , and Ca II IRT emission for both stars is mostly non-variable, with only a few spectra showing excess emission that is attributable to flaring events or prominences. In particular, TWA 6 shows excess red-shifted emission in the H α , H β , and Ca II IRT lines in three spectra; however, these features are not long lasting and are not periodic. Indeed, the magnetic topology at these phases is such that excess emission could be due to off-limb prominence material that is rotating away from the observer in closed magnetic loops.

Using ZDI, we have derived a surface brightness map of TWA 6 and the magnetic topologies of both stars. We find that TWA 6 has many cool spots and warm plages on its surface, with a total coverage of around 17 per cent. We detect no significant modulation of the Stokes I lines profiles for TWA 8A, and so find its surface to be compatible with a uniformly bright star. The reconstructed magnetic fields for TWA 6 and TWA 8A are somewhat different in strength, and dramatically different in topology. TWA 6 has a field that is split equally between poloidal and toroidal components, with the largest fraction of energy in higher order modes (with $\ell > 3$), with a total unsigned flux of $\langle B \rangle = 840 \text{ G}$ and where the large-scale magnetosphere is tilted at 35° from the rotation axis. On the other hand, TWA 8A has a highly poloidal field, with most of the energy in the high-order modes with $\ell > 3$. The field strength is sufficiently large that the Stokes I lines profiles are significantly Zeeman broadened, with Zeeman signatures clearly detected in individual Stokes V spectral lines. We derive a total unsigned flux of $\langle B \rangle = 1.4 \text{ kG}$, using a magnetic filling factor f equal to 0.2 (meaning that 20 per cent of the surface was covered with the mapped magnetic features), where on large scales the magnetosphere is tilted at 20° from the rotation axis.

For TWA 8A, our simultaneous fits to both Stokes I and V spectra yields a mean magnetic field strength of $\langle |B_f| \rangle = 5.9 \pm 0.2 \text{ kG}$, with a significant fraction of energy in high-strength fields ($> 5 \text{ kG}$). Given that we recover a larger fraction of high magnetic field strengths

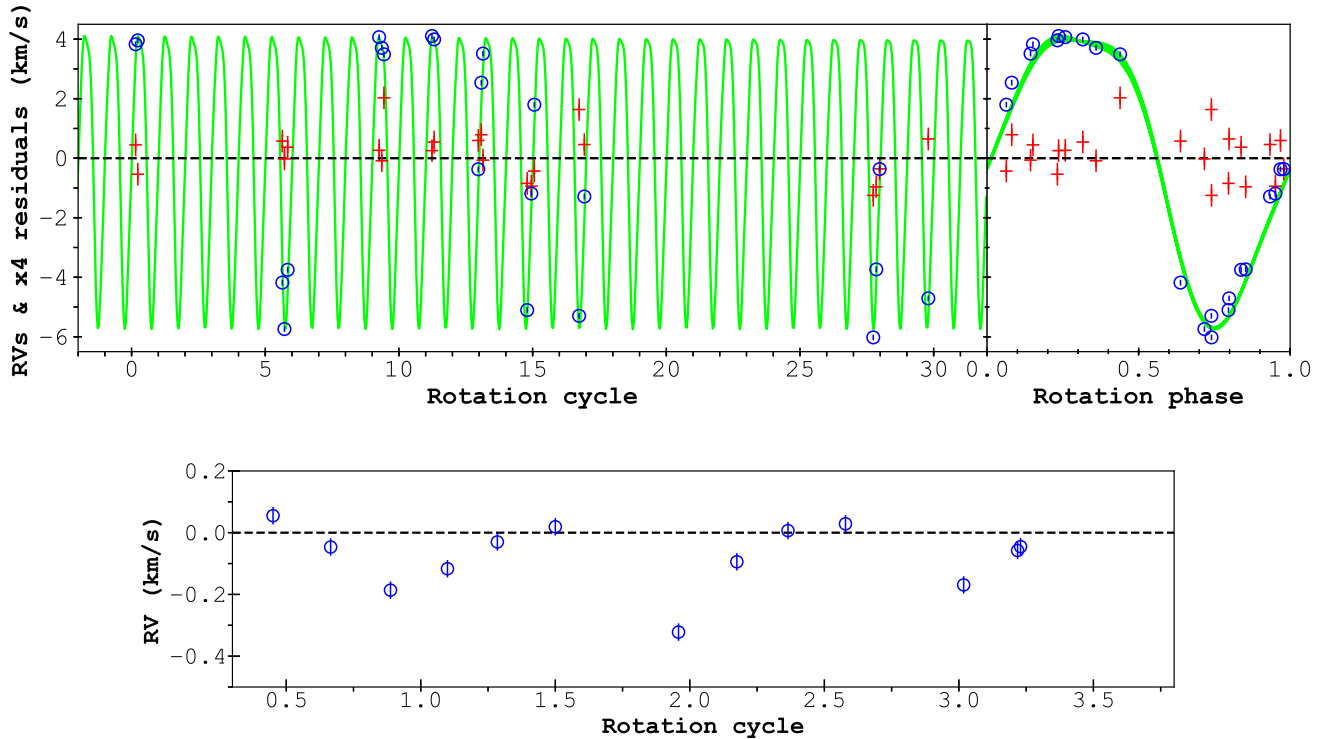


Figure 12. Top left-hand panel: RV variations (in the stellar rest frame) of TWA 6 a function of rotation phase, as measured from our observations (open blue circles) and predicted by the tomographic brightness map of Fig. 6 (green line). RV residuals are also shown (red crosses, with values and error bars scaled by a factor of 4 for clarity), and exhibit a rms dispersion equal to 0.20 km s^{-1} . RVs are estimated as the first-order moment of the Stokes I LSD profiles rather than through Gaussian fits, due to their asymmetric and often irregular shape. Top right-hand panel: Same as the top left-hand panel after phase-folding the data and model. Note that the model shows little variation over the ~ 30 rotation cycles, showing the very low level of differential rotation. Bottom panel: The measured RVs of TWA 8A as a function of rotation phase. Note that the filtered RVs are not shown for TWA 8A as the line profiles are compatible with a star of uniform brightness. The unfiltered RVs show a period signal that is equal to the stellar rotation period. This figure is best viewed in colour.

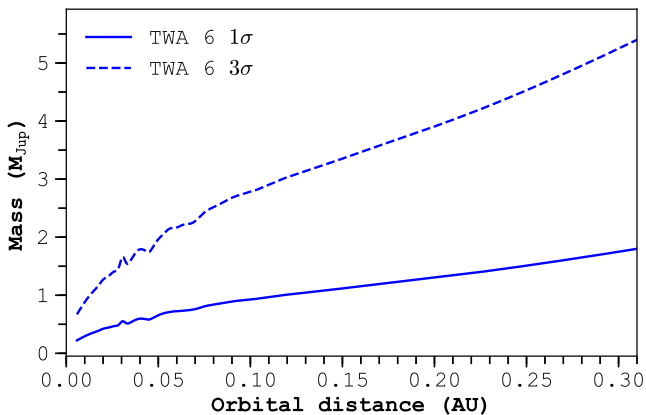


Figure 13. The 1σ and 3σ upper limits (solid and dashed lines, respectively) on the recovered planet mass as a function of orbital distance, using the RVs shown in Fig. 12 for TWA 6. This figure is best viewed in colour.

from our direct modelling of Stokes I profiles, with those fields having small filling factors, a significant proportion of magnetic energy likely lies in small-scale fields that are unresolved by ZDI. The difference between direct spectral fitting and ZDI is likely due to several factors. First, by the cancellation of near-by regions of different sign in Stokes V (providing most of the difference between Stokes I and V in single lines); Secondly, by the signal in Stokes V LSD profiles not adding perfectly coherently due to the non-self-similarity of different lines in Stokes V , with scaling amplitudes by

effective Landé factors yielding a less accurate line profile (most of the difference between single lines and LSD profiles). Hence, small-scale high-strength magnetic fields are not recovered with LSD and are thus not reconstructed with ZDI.

Compared to Tap 26, another wTTS that has a similar mass, age and rotation rate (Yu et al. 2017), TWA 6 has a larger toroidal field component (50 per cent for TWA 6 versus 30 per cent for Tap 26), with a total field strength that is around twice as large. Likewise, the field of TWA 6 is also around twice as strong as those of the slower rotating (but similarly massive) wTTSs, V819 Tau and V830 Tau (Donati et al. 2015). In the case of TWA 8A, we find that it has a weaker (poloidal) dipole field (of $B = 0.72 \text{ kG}$) compared to LkCa 4 (with $B = 1.6 \text{ kG}$), a wTTSs with a similar rotation rate and a slightly higher mass ($P_{\text{rot}} = 3.374 \text{ d}$, $0.8 M_{\odot}$). Moreover, compared to main-sequence M dwarfs with a similar mass and period, namely EV Lac ($\langle B \rangle = 0.57 \text{ kG}$) and GJ 182 ($\langle B \rangle = 172 \text{ G}$), we see that TWA 8A has a slightly stronger magnetic field.

In Fig. 14, we compare the magnetic field topologies of all cTTSs and wTTSs so far mapped with ZDI in an H–R diagram. Fig. 14 also indicates the fraction of the field that is poloidal, the axisymmetry of the poloidal component, and shows PMS evolutionary tracks from Siess et al. (2000). In contrast to cTTSs of the MaPP project, the wTTSs that have been analysed (so far) in the MaTY SSE sample do not appear to show many obvious trends with internal structure. The magnetic field strength does not appear to change significantly after the star becomes mostly radiative, with the largely convective V830 Tau, V819 Tau and V410 Tau hosting a similarly strong dipole field to the mostly radiative TAP 26, and with the

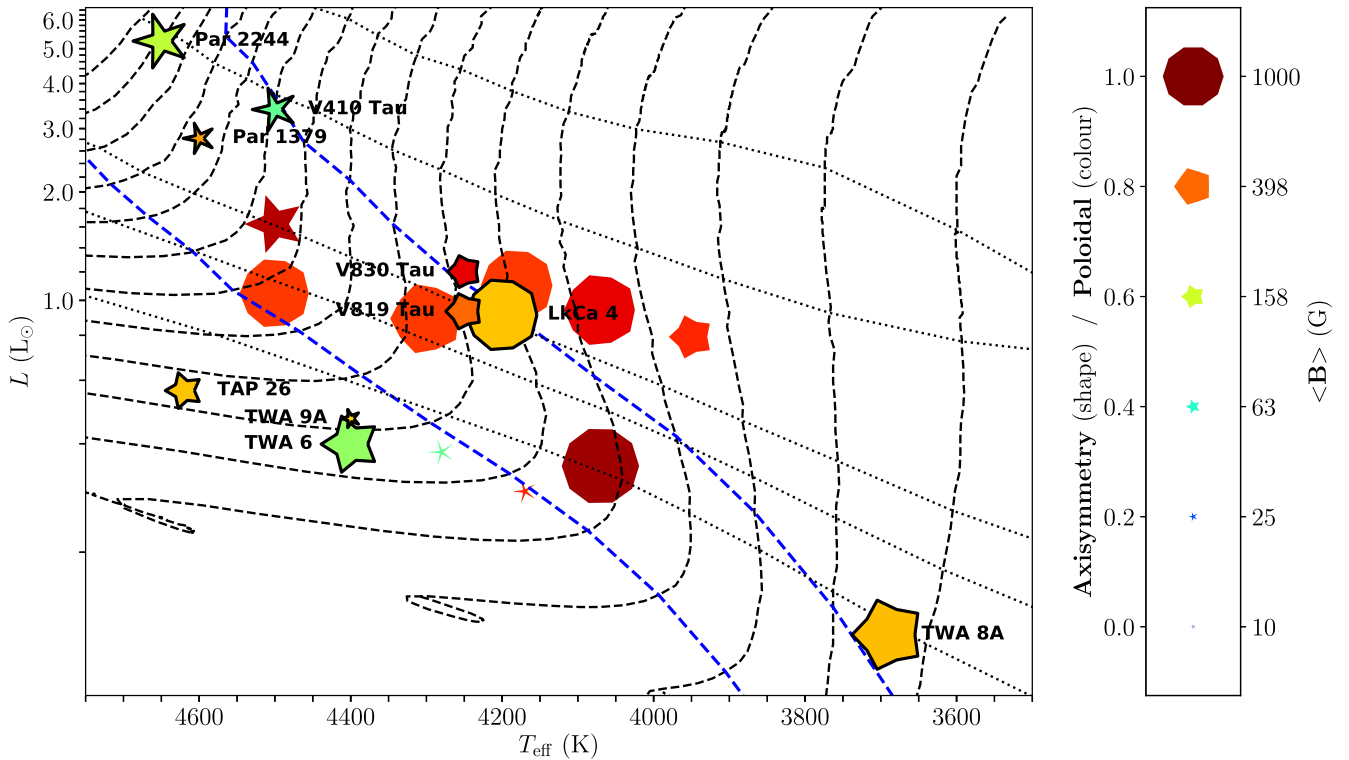


Figure 14. H–R diagram showing the MaTYSSSE wTTs (black line border and labelled) and the MaPP cTTs (no border). The size of the symbols represents the surface-averaged magnetic field strength (with a larger symbol meaning a stronger field), the colour of the symbol represents the fraction of the field that is poloidal (with red being completely poloidal), and the shape of the symbols represents the axisymmetry of the poloidal field component (with higher axisymmetry shown as a more circular symbol). Also shown are evolutionary tracks from Siess et al. (2000) (black dashed lines, ranging from 0.3 to 1.9 M_{\odot}), with corresponding isochrones (black dotted lines, for ages of 0.5, 1, 3, 5, and 10 Myr), and lines showing 100 per cent and 50 per cent convective interior by radius (blue dashed).

largely convective Par 2244 hosting a similarly strong field mostly radiative TWA 6. Moreover, the percentage of poloidal field does not appear to change from when the star is fully convective to when it is mostly radiative (e.g. V410 Tau and TWA 6 are both around 50 per cent poloidal). However, the degree of axisymmetry of the poloidal field appears to correlate with the strength of the magnetic field, given that LkCa 4 and TWA 8A (two stars with significantly stronger fields of 1.2 kG and 1.4 kG, respectively) are mostly axisymmetric ($\gtrsim 70$ per cent). Considering both cTTs and wTTs as a whole, it appears that stars are mostly poloidal and axisymmetric when they are mostly convective and cooler than ~ 4300 K. Moreover, stars hotter than ~ 4300 K appear to be less axisymmetric and less poloidal, regardless of their internal structure. We note that the wTTs studied thus far clearly show a wider range of field topologies compared to those of cTTs, with large-scale fields that can be more toroidal and non-axisymmetric, consistent with the fact that most of them are largely radiative or are higher mass. We also note that a more complete analysis will be possible once the remainder of the MaTYSSSE sample has been analysed.

Through our tomographic modelling, we were able to determine that TWA 6 has a non-zero surface latitudinal-shear at a confidence level of over 99.99 per cent for the brightness map, and 90 per cent for the magnetic map, as measured over the 16 nights of observation. Its shear rate is around 56 times smaller than the Sun, with an equator-pole lap time of 640^{+110}_{-80} d. Given the lack of variability in the lines profiles and the small number of observed rotations (~ 3 cycles), we were unable to measure the shear rate

for TWA 8A. Out measured shear rate for TWA 6 is similar to that found for V410 Tau, V819 Tau, V830 Tau, and LkCa 4 (Skelly et al. 2010; Donati et al. 2014, 2015), which are all of similar mass.

Finally, the brightness map of TWA 6 was used to predict the activity-related RV jitter due to stellar activity, allowing us to filter the measured RVs in the search for potential hJs (in the same manner as Donati et al. 2014, 2015). Here, the activity jitter was filtered down to a rms RV precision of ~ 0.20 km s $^{-1}$ (from an initial unfiltered rms of 3.8 km s $^{-1}$). While this is well above the RV precision of ESPaDOnS, the high $v \sin i$ decreases the accuracy of the filtering process, with an intrinsic uncertainty of around 0.1 km s $^{-1}$. We find no significant peaks in a periodogram analysis, and find that TWA 6 is unlikely to host a hJ with an orbital period in the range of what we can detect, with a 3σ error bar on the semi-amplitude of the RV residuals equal to 0.19 km s $^{-1}$, translating into a planet mass of $\simeq 3.1 M_{\text{Jup}}$ orbiting at $\simeq 0.1$ au.

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APPENDIX A: LINE PROFILES OF CA II IR TRIPLET, H α , AND H β FOR TWA 6 AND TWA 8A

Line profiles of the Ca II IR triplet, H α , and H β are shown in Figs A1 and A2 for TWA 6 and TWA 8A, respectively. Fig. A3 shows Stokes V line profiles of the Ca II IR triplet and H α , for TWA 8A. Figs A4 and A5 show normalized variance profiles and autocorrelation matrices of H α , H β , and He I D3, for TWA 8A and TWA 6, respectively.

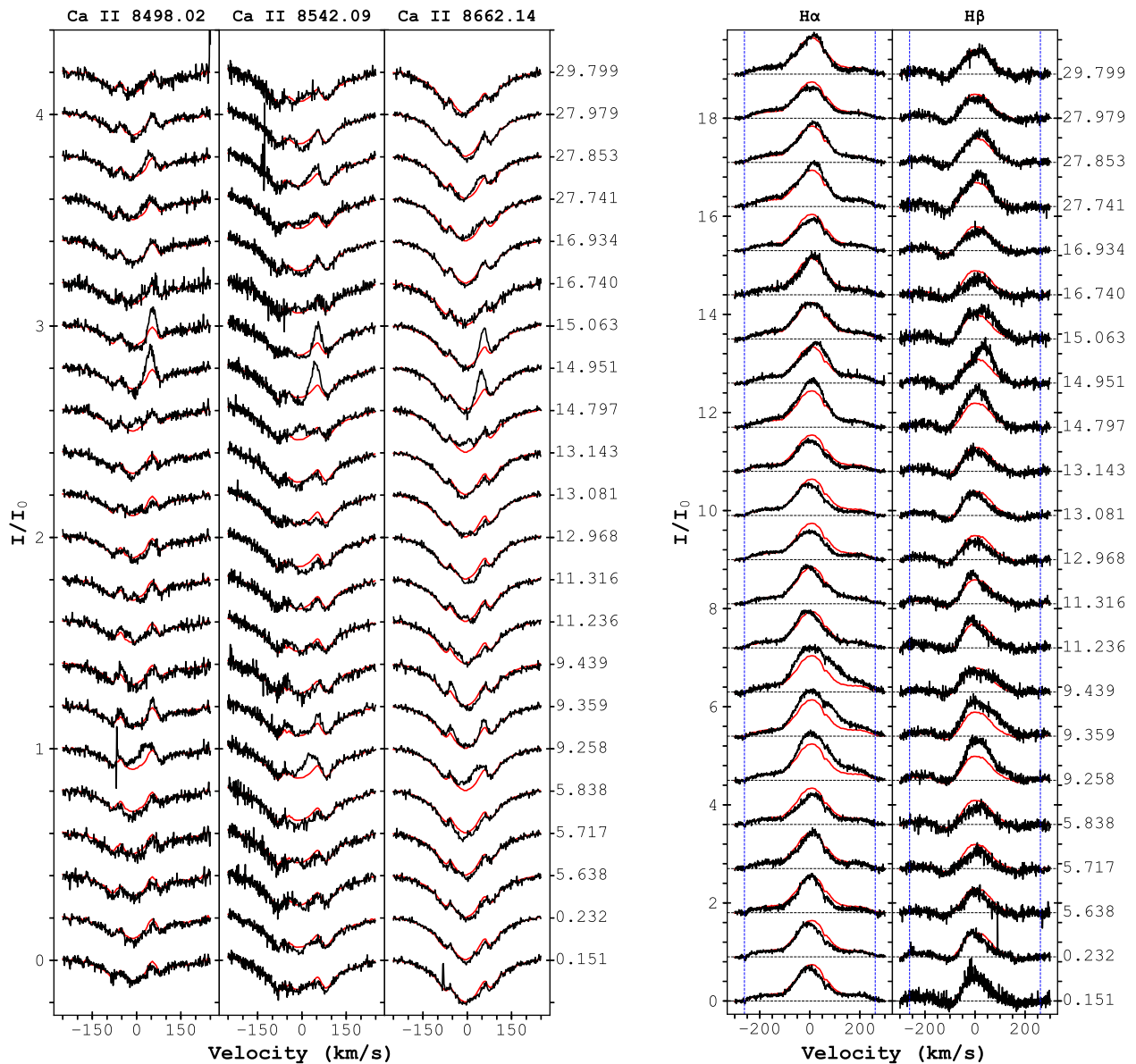


Figure A1. For TWA 6. Left-hand panel: The Ca II IR triplet, with line profiles of the 8498.02, 8542.09, and 8662.14 Å components shown (left to right) as black solid lines, where the mean line profile is shown in red, with the cycle number displayed on the right of the profiles. Right-hand panel: H α and H β line profiles, shown in the same manner, additionally showing the co-rotation radius as a dashed blue line.

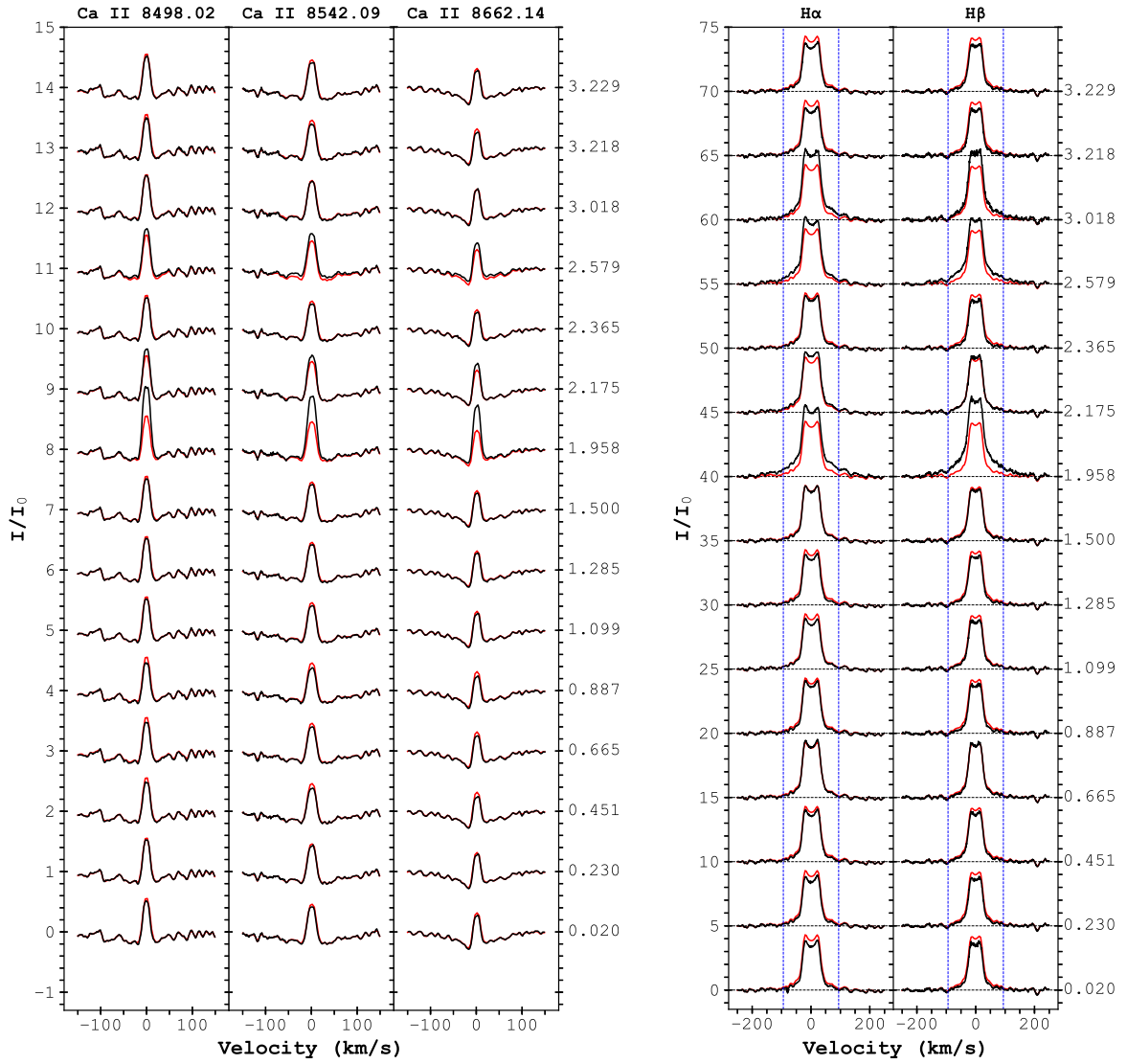


Figure A2. Same as Fig. A1 but for TWA 8A, with the left-hand panel showing the Ca II IR triplet, and the right-hand panel showing H α and H β line profiles.

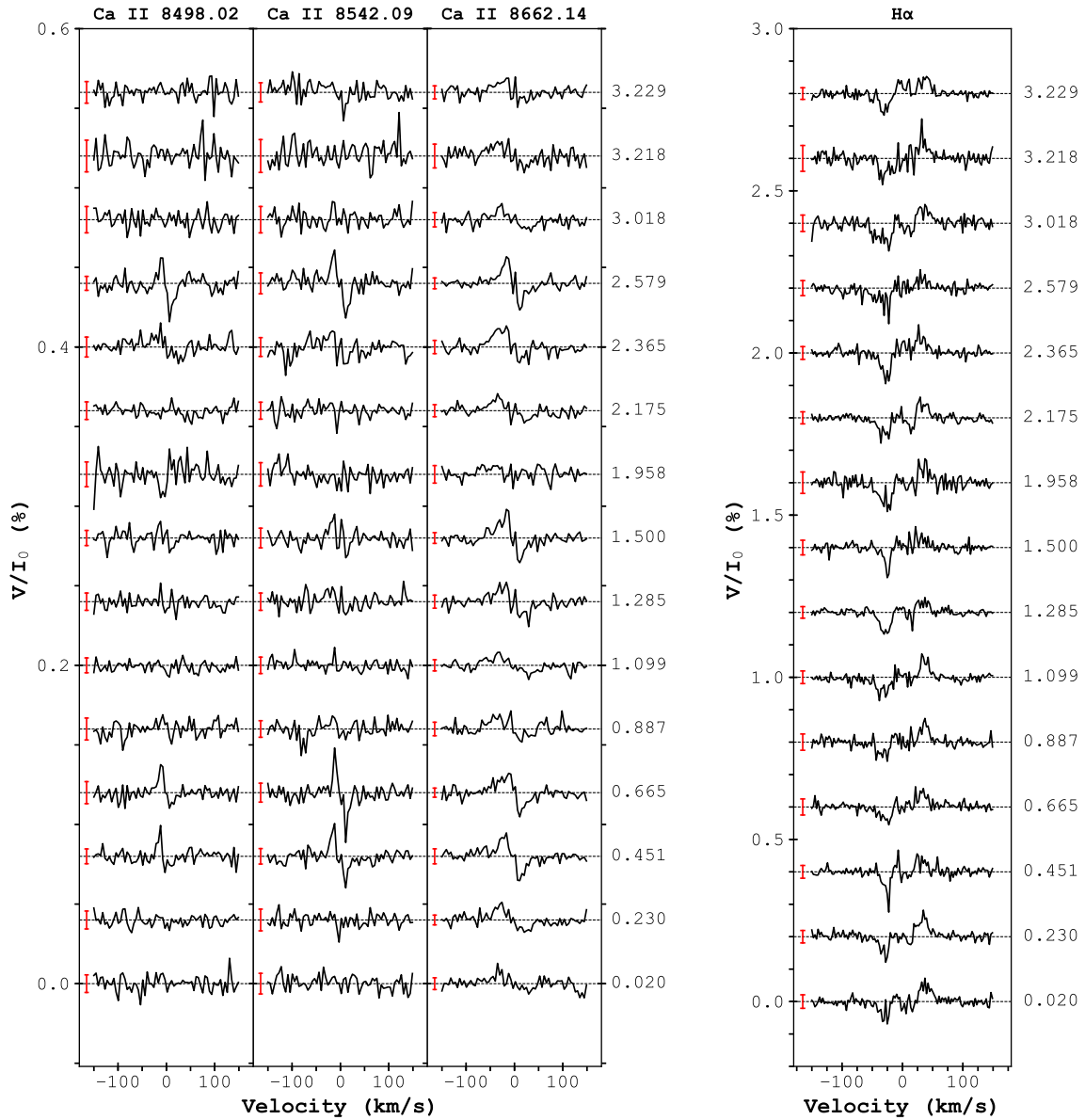


Figure A3. Stokes V line profiles of TWA 8A with the Ca II IRT shown in the left-hand panel, and H α shown in the right-hand panel. 3σ error bars are shown in red on the left-hand side of the line profiles.

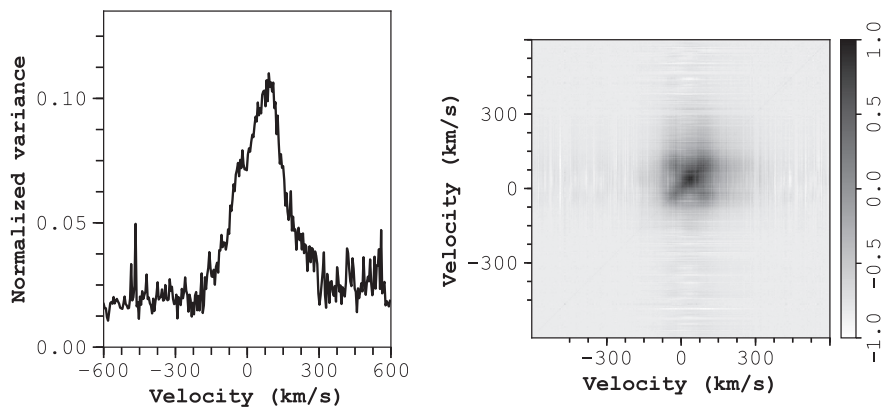


Figure A4. Left-hand panel: The normalized variance profile of H α for TWA 6. There is variance from around -200 km s^{-1} up to around $+300 \text{ km s}^{-1}$. Right-hand panel: The autocorrelation matrix for H α , where black means perfect correlation and white means perfect anticorrelation.

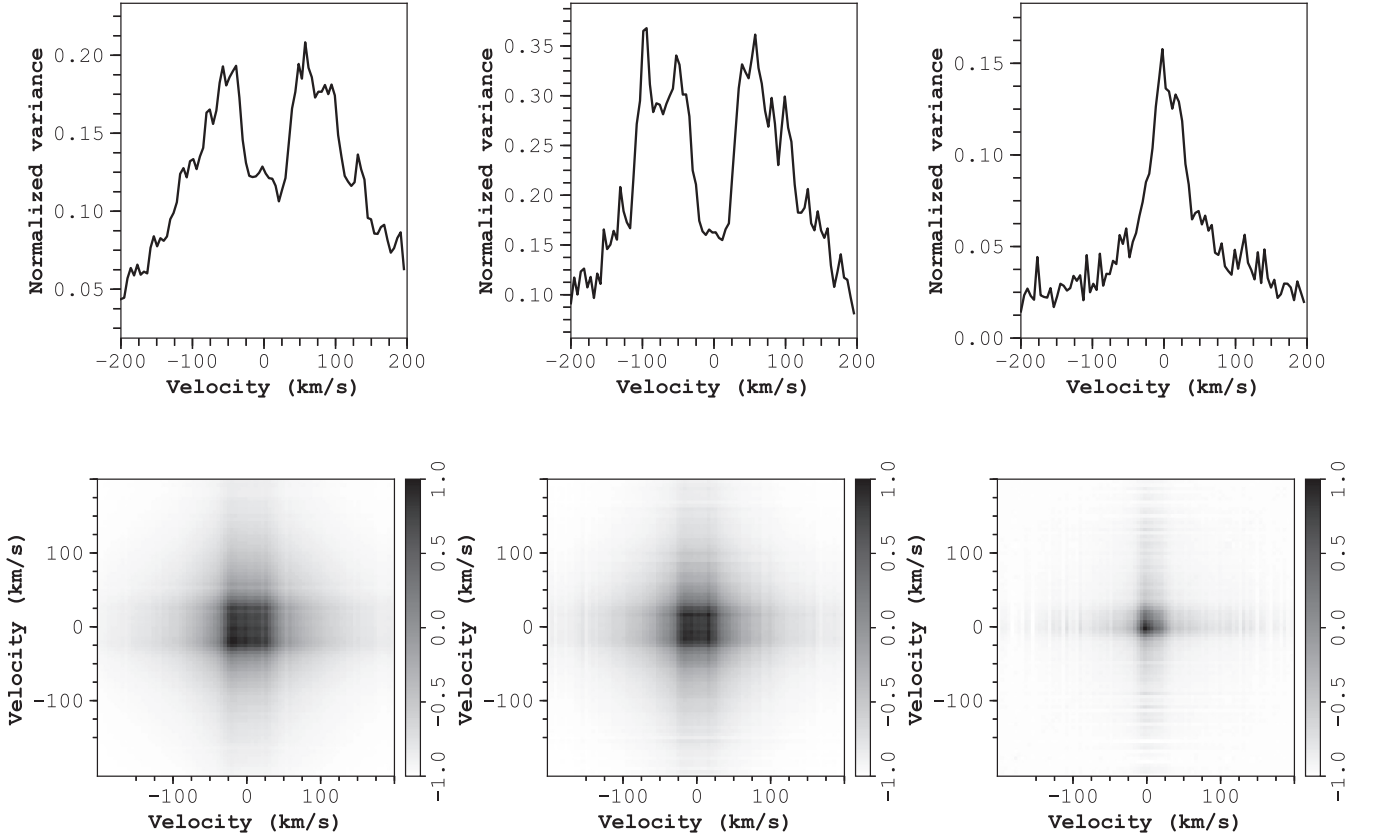


Figure A5. Same as Fig. A4 but for TWA 8A. The top row shows normalized variance profiles for H α , H β , and He I D3 (left to right), with the bottom row showing the corresponding autocorrelation matrices.

APPENDIX B: ACCRETION STATUS OF TWA 6 AND TWA 8A

The SEDs of TWA 6 and TWA 8A show no evidence of an IR excess (see Fig. 2), suggesting that both stars are discless. Nevertheless, we may use our high-quality spectra of both targets to determine their accretion status using several metrics.

Following our previous studies (e.g. Hill et al. 2017), one may estimate the level of surface accretion in TTSs by adopting the relations between line luminosity L_{line} and the accretion luminosity L_{acc} of Alcalá et al. (2017). For this purpose, we determined L_{line} by assuming blackbody scaling using the stellar radius R_{\star} and T_{eff} given in Table 2. Then, the mass accretion rate \dot{M}_{acc} was calculated using the relationship

$$\dot{M}_{\text{acc}} = \frac{L_{\text{acc}} R_{\star}}{GM_{\star} \left(1 - \frac{R_{\star}}{R_{\text{in}}}\right)}, \quad (\text{B1})$$

where R_{in} denotes the truncation radius of the disc, and is taken to be $5R_{\star}$ (Gullbring et al. 1998).

For TWA 6, we detect weak He I D₃ emission with an EW of around 0.03 Å (1.6 km s⁻¹), corresponding to $\log \dot{M}_{\text{acc}} \simeq -10.6 M_{\odot} \text{ yr}^{-1}$. We find the H β emission (see Fig. A1) to have an EW ranging between 0.69 and 1.44 Å (average of 0.96 Å, equivalent to 59.5 km s⁻¹, corresponding to $\log \dot{M}_{\text{acc}} \simeq -10.1 M_{\odot} \text{ yr}^{-1}$), and the H α emission to have an EW ranging between 2.26 and 4.03 Å (average of 2.85 Å, equivalent to 130.2 km s⁻¹, corresponding to $\log \dot{M}_{\text{acc}} \simeq -10.2 M_{\odot} \text{ yr}^{-1}$).

For TWA 8A, we detect moderate He I D₃ emission with an EW of around 0.3 Å (17.4 km s⁻¹), corresponding to

$\log \dot{M}_{\text{acc}} \simeq -9.7 M_{\odot} \text{ yr}^{-1}$. We find the H β emission (see Fig. A2) to have an EW ranging between 3.1 and 6.8 Å (average of 3.9 Å, equivalent to 238.0 km s⁻¹, corresponding to $\log \dot{M}_{\text{acc}} \simeq -9.7 M_{\odot} \text{ yr}^{-1}$), and the H α emission to have an EW ranging between 5.9 and 10.9 Å (average of 7.2 Å, equivalent to 326.9 km s⁻¹, corresponding to $\log \dot{M}_{\text{acc}} \simeq -9.8 M_{\odot} \text{ yr}^{-1}$).

These accretion rates would suggest that both stars are weakly accreting; however, as discussed in Hill et al. (2017), chromospheric activity in TTS becomes a significant influence on the strength and width of emission lines in the low accretion regime. Pertinently, the large convective turnover times of TTSs (Gilliland 1986) combined with their rapid rotation means they possess a low Rossby number, placing them well within the saturated activity regime (e.g. Reiners, Schüssler & Passetger 2014). Indeed, the H α line luminosity is observed to saturate in young stars at around $\log [L(\text{H } \alpha)/L_{\text{bol}}] = -3.3$ or lower (Barrado y Navascués & Martín 2003), and as both our target stars show line luminosities that are similar to (or below) this level, with $\log [L(\text{H } \alpha)/L_{\text{bol}}]$ equal to -3.19 for TWA 6 and -3.78 for TWA 8A, any estimate of accretion rates based on line luminosities (especially H α , and to a lesser extent H β) must be considered to be significantly influenced or even dominated by chromospheric activity.

The distinction between emission due to accretion and that due to chromospheric activity has been characterized by several authors, yielding a distinct threshold between these regimes. Using the empirical spectral-type-dependant relationship between the EW(H α) and the accretion rate of Barrado y Navascués & Martín (2003), the defining threshold of an accreting TTS is EW(H α) equal to 5.1 and 12.2 Å for K5 and M3 spectral types, respectively

(appropriate for TWA 6 and TWA 8A, see Section 3.1). Given that the maximum $\text{EW}(\text{H}\alpha)$ of TWA 6 and TWA 8A are equal to 4.0 and 10.9 Å, both stars lie below these limits and fall into the non-accreting regime (where line broadening is dominated by chromospheric activity).

Elsewhere, Manara et al. (2017) derived an empirical relationship between a star’s spectral-type and the point at which line emission may be dominated by chromospheric activity (termed chromospheric accretion ‘noise’). In the case of TWA 6, this threshold is at $\log(L_{\text{acc, noise}}/L_{\text{star}}) = -2.3 \pm 0.1$. Given that the average line luminosities $\log(L_{\text{acc}}/L_{\text{star}})$ for $\text{H}\alpha$, $\text{H}\beta$, and He I D_3 are, respectively, equal to -3.19 ± 0.01 , -3.04 ± 0.02 , and -3.63 ± 0.03 , the luminosity of all three emission lines are significantly below the threshold of chromospheric noise. Likewise for TWA 8A, this threshold is estimated as $\log(L_{\text{acc, noise}}/L_{\text{star}}) = -2.59 \pm 0.13$. Here, the average line luminosities for $\text{H}\alpha$, $\text{H}\beta$, and

He I D_3 are, respectively, equal to -3.78 ± 0.02 , -3.65 ± 0.03 , and -3.62 ± 0.02 , where again, all emission is well below the threshold where one can distinguish between accretion and chromospheric emission.

Thus, the accretion rates determined above for TWA 6 and TWA 8A must be taken to be upper limits, given that chromospheric emission is likely the dominant broadening mechanism. Hence, our target stars are likely not accreting (or are doing so at an undetectable level), thus confirming their classification as wTTSs – a result consistent with past work by White & Hillenbrand (2004) and Kastner et al. (2016).

APPENDIX C: MAGNETIC FIELDS FROM DIRECT SPECTRAL FITTING

See Tables C1 and C2.

Table C1. Atomic data used in the direct spectrum fitting, from VALD, for the major lines. Additional much weaker lines were included in the spectrum synthesis for completeness, but are omitted here for brevity. The quantities J_{low} and J_{high} refer to the lower and upper levels of the transition, respectively. Term symbols are provided to identify lines of the same multiplet.

Species	Wavelength (Å)	$\log gf$	$E_{\text{low}} (E_v)$	J_{low}	J_{high}	Landé g_{low}	Landé g_{high}	Multiplet terms
Ti sc i	9675.54	-0.804	0.8360	4	4	1.34	1.35	$a^5F - z^5F^\circ$
Ti sc i	9688.87	-1.610	0.8129	1	2	1.00	1.50	$a^5F - z^5F^\circ$
Ti sc i	9705.66	-1.009	0.8259	3	3	1.26	1.26	$a^5F - z^5F^\circ$
Ti sc i	9718.96	-1.181	1.5025	4	3	1.00	0.95	$a^1G - z^1F^\circ$
Ti sc i	9728.41	-1.206	0.8181	2	2	1.00	1.00	$a^5F - z^5F^\circ$
Ti sc i	9743.61	-1.306	0.8129	1	1	0.00	0.00	$a^5F - z^5F^\circ$
Ti sc i	9770.30	-1.581	0.8484	5	4	1.34	1.55	$a^5F - z^5F^\circ$
Ti sc i	9783.31	-1.428	0.8360	4	3	1.26	1.48	$a^5F - z^5F^\circ$
Ti sc i	9783.59	-1.617	0.8181	2	1	0.00	1.49	$a^5F - z^5F^\circ$
Ti sc i	9787.69	-1.444	0.8259	3	2	1.00	1.50	$a^5F - z^5F^\circ$
Ti sc i	9832.14	-1.130	1.8871	5	4	1.21	1.21	$a^3G - y^3F^\circ$

Table C2. Best-fitting parameters from direct spectral fitting of TWA 8A, using our third model that fits Stokes I and V simultaneously, for observations that could be adequately telluric corrected. Each column gives the fitted parameters for the spectrum obtained on the date given at the top. Mean values are presented in Table 3.

	2015-03-25	2015-03-26	2015-03-27	2015-03-28	2015-03-29	2015-03-30	2015-03-31	2015-04-01	2015-04-05	2015-04-06
$u \sin i$ (km s $^{-1}$)	4.75 ± 0.19	4.78 ± 0.17	4.72 ± 0.18	5.08 ± 0.17	4.61 ± 0.18	4.66 ± 0.16	4.66 ± 0.16	4.95 ± 0.17	4.92 ± 0.17	5.02 ± 0.17
v_{mic} (km s $^{-1}$)	1.18 ± 0.07	1.07 ± 0.06	1.14 ± 0.06	1.13 ± 0.06	1.08 ± 0.06	1.06 ± 0.05	1.06 ± 0.06	1.08 ± 0.06	1.01 ± 0.06	1.02 ± 0.06
[Ti/H]	-7.007 ± 0.014	-6.985 ± 0.012	-7.005 ± 0.013	-6.974 ± 0.012	-6.981 ± 0.012	6.988 ± 0.011	-6.966 ± 0.011	-6.963 ± 0.012	-6.934 ± 0.011	-6.952 ± 0.012
+2 kG	0.168 ± 0.023	0.182 ± 0.017	0.167 ± 0.018	0.154 ± 0.018	0.150 ± 0.018	0.176 ± 0.016	0.175 ± 0.016	0.165 ± 0.018	0.147 ± 0.019	0.125 ± 0.019
+5 kG	0.229 ± 0.010	0.227 ± 0.009	0.259 ± 0.010	0.239 ± 0.010	0.221 ± 0.010	0.228 ± 0.009	0.239 ± 0.009	0.268 ± 0.010	0.258 ± 0.010	0.282 ± 0.010
+10 kG	0.053 ± 0.008	0.061 ± 0.006	0.051 ± 0.007	0.056 ± 0.007	0.052 ± 0.006	0.056 ± 0.006	0.061 ± 0.006	0.055 ± 0.006	0.058 ± 0.006	0.048 ± 0.007
+15 kG	0.052 ± 0.007	0.044 ± 0.006	0.049 ± 0.006	0.051 ± 0.006	0.045 ± 0.006	0.047 ± 0.005	0.043 ± 0.005	0.043 ± 0.006	0.045 ± 0.006	0.047 ± 0.006
+20 kG	0.034 ± 0.005	0.030 ± 0.004	0.027 ± 0.005	0.042 ± 0.005	0.038 ± 0.005	0.034 ± 0.004	0.037 ± 0.004	0.037 ± 0.005	0.038 ± 0.004	0.038 ± 0.005
-2 kG	0.163 ± 0.023	0.166 ± 0.017	0.145 ± 0.018	0.159 ± 0.018	0.148 ± 0.018	0.173 ± 0.016	0.161 ± 0.016	0.144 ± 0.018	0.146 ± 0.019	0.141 ± 0.019
-5 kG	0.221 ± 0.010	0.202 ± 0.009	0.205 ± 0.010	0.202 ± 0.010	0.217 ± 0.010	0.213 ± 0.009	0.199 ± 0.009	0.209 ± 0.010	0.205 ± 0.010	0.213 ± 0.010
-10 kG	0.044 ± 0.008	0.045 ± 0.006	0.036 ± 0.007	0.053 ± 0.007	0.041 ± 0.006	0.041 ± 0.006	0.036 ± 0.006	0.039 ± 0.006	0.035 ± 0.006	0.040 ± 0.007
-15 kG	0.010 ± 0.006	0.000 ± 0.006	0.010 ± 0.006	0.010 ± 0.006	0.012 ± 0.006	0.010 ± 0.005	0.012 ± 0.005	0.011 ± 0.006	0.009 ± 0.006	0.013 ± 0.006
-20 kG	0.022 ± 0.005	0.028 ± 0.004	0.024 ± 0.005	0.035 ± 0.005	0.023 ± 0.005	0.023 ± 0.004	0.024 ± 0.004	0.028 ± 0.005	0.031 ± 0.004	0.029 ± 0.005
0 kG	0.004 ± 0.039	0.017 ± 0.031	0.028 ± 0.033	0.000 ± 0.033	0.051 ± 0.032	0.000 ± 0.029	0.014 ± 0.029	0.000 ± 0.032	0.028 ± 0.033	0.023 ± 0.033

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