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LETTER TO THE EDITOR

# Detection of CS in Neptune's atmosphere from ALMA observations

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## ABSTRACT

**Context.** The large and vertically non-uniform abundance of CO in Neptune's atmosphere has been interpreted as the result of past cometary impact(s), either single or distributed in size and time, which could also be at the origin of Neptune's HCN.

**Aims.** We aim to provide observational support for this scenario by searching for other comet-induced species, in particular carbon sulfide (CS) which has been observed continuously in Jupiter since the 1994 Shoemaker-Levy 9 impacts.

**Methods.** In April 2016 we used the ALMA interferometer to search for CS(7-6) at 342.883 GHz in Neptune.

**Results.** We report on the detection of CS in Neptune's atmosphere, the first unambiguous observation of a sulfur-bearing species in a giant planet beyond Jupiter. Carbon sulfide appears to be present only at submillibar levels, with a column density of  $(2.0\text{--}3.1) \times 10^{12} \text{ cm}^{-2}$ , and a typical mixing ratio of  $(2\text{--}20) \times 10^{-11}$  that depends on its precise vertical location. The favoured origin of CS is deposition by a putative large comet impact several centuries ago, and the strong depletion of CS with respect to CO – compared to the Jupiter case – is likely due to the CS sticking to aerosols or clustering to form polymers in Neptune's lower stratosphere.

**Conclusions.** The CS detection, along with recent analyses of the CO profile, reinforces the presumption of a large comet impact into Neptune ~1000 yr ago, that delivered CO, CS, and HCN at the same time.

**Key words.** planets and satellites: individual: Neptune – techniques: spectroscopic – submillimeter: planetary systems

## 1. Introduction

External sources of oxygen are ubiquitous in the giant planets and Titan. This is demonstrated from the presence, in all five objects, of condensable species such as H<sub>2</sub>O and CO<sub>2</sub> at stratospheric levels (Feuchtgruber et al. 1997; Lellouch et al. 2002; Orton et al. 2014), and from non-vertically uniform distributions of CO in all four giant planets (Bézar et al. 2002; Cavalié et al. 2009, 2014; Lellouch et al. 2005; Hesman et al. 2007; Luszcz-Cook & de Pater 2013). The CO abundance in Neptune's stratosphere (about 1 ppm) is by far the largest of all the giant planets, and the external CO component implies a time-averaged influx of oxygen material of the order of  $10^8 \text{ cm}^{-2} \text{ s}^{-1}$  (Lellouch et al. 2005; Luszcz-Cook & de Pater 2013), at least 3 orders of magnitude larger than for Uranus. Building on the similarity with Jupiter, where the Shoemaker-Levy 9 (SL9) impacts injected massive amounts of fresh CO that added to the pre-existing levels on the planet, the preferred explanation for Neptune's CO is delivery from past cometary impacts in the form of a single large impact ~200 yr ago (Lellouch et al. 2005; Hesman et al. 2007) or of a distribution of impactors in size and time (Luszcz-Cook & de Pater 2013). A comet impact would also supply HCN (Lellouch et al. 2005), although other sources of Neptune's HCN are possible (Marten et al. 1993; Lellouch et al. 1994).

The SL9 impacts also generated numerous sulfur species in Jupiter's stratosphere, including S<sub>2</sub>, H<sub>2</sub>S, CS, CS<sub>2</sub>, and OCS (Noll et al. 1995; Lellouch et al. 1995). These species, except CS and CS<sub>2</sub>, have short atmospheric lifetimes (Moses 1996). In contrast, CS has been continuously observed to be present in Jupiter's atmosphere over the last 20 yr, with a relatively slow decay rate. This makes CS an attractive target in the framework of the comet impact scenario for Neptune. We report here the first detection of CS in Neptune's atmosphere, using ALMA.

## 2. Observations and data reduction

We observed Neptune with ALMA on three occasions between April 23 and 30, 2016. At that time the array included 39–41 12 m antennas in configuration C36-2/3, providing a synthetic beam of about 0.6". The spectral set-up included the CS(7-6) line at 342.883 GHz<sup>1</sup> at a spectral resolution of 1 MHz. The CO(3-2) and HCN(4-3) lines were also covered. Observations (see details in Table 1) used the strong quasar J0006-0623 to calibrate the RF bandpass and the quasar J2246-1206 as phase calibrator. Initial data reduction was performed under the CASA software (i.e. the ALMA data reduction pipeline<sup>2</sup>).

The resulting calibrated visibilities were then exported into the GILDAS package<sup>3</sup> for i) an improved amplitude calibration by scaling the three datasets to a single reference Neptune's angular diameter of 2.24"; ii) a self-calibration technique using Neptune's continuum, which was also used as the absolute flux reference, yielding to an accuracy of 5%; and iii) a spectral extraction by fitting the data directly in the UV-plane (i.e. the Fourier plane, as opposed to the image plane), providing the disk-averaged flux with optimum signal-to-noise ratio (S/N). These steps resulted in the detection of CS in three independent datasets (Fig. 1). The averaged line, which shows a contrast of 55 mJy (i.e. 0.22% of Neptune's continuum) with a S/N of 7 and a linewidth of 5.2 MHz was used for modelling.

## 3. Modelling and results

Data were modelled with a standard radiative transfer model fully accounting for spherical geometry and broadening due to planetary rotation, developed for Jupiter (Moreno et al. 2001)

<sup>1</sup> Two slightly different frequency setups were used on April 23–30.

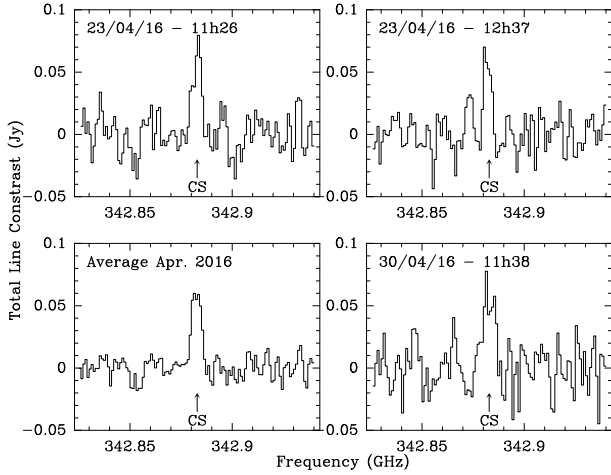
<sup>2</sup> <https://casa.nrao.edu/>

<sup>3</sup> <http://www.iram.fr/IRAMFR/GILDAS>

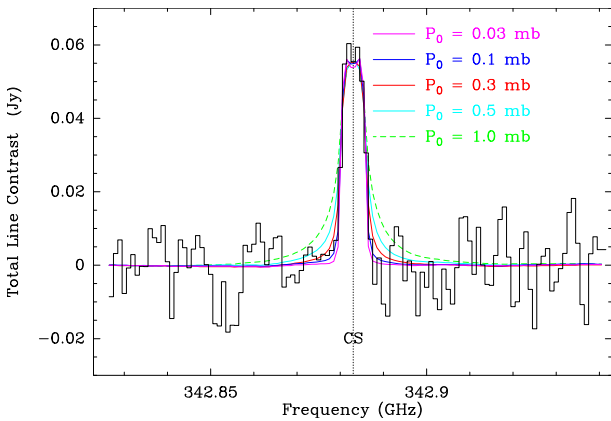
**Table 1.** ALMA observations of Neptune.

UT start date yyyy mm dd.ddd	$\Delta^1$ [AU]	$\theta^2$ [ $''$ ]	$A^3$	$T_{\text{int}}^4$ (min)	PWV <sup>5</sup> (mm)
2016 04 23.477	30.559	2.234	41	20	0.9
2016 04 23.525	30.559	2.234	39	20	0.9
2016 04 30.484	30.462	2.241	41	20	1.4

**Notes.** (1) Geocentric distance. (2) Apparent diameter at  $p = 1$  bar. (3) Number of active antennas. (4) On-source integration time. (5) Precipitable water vapour.

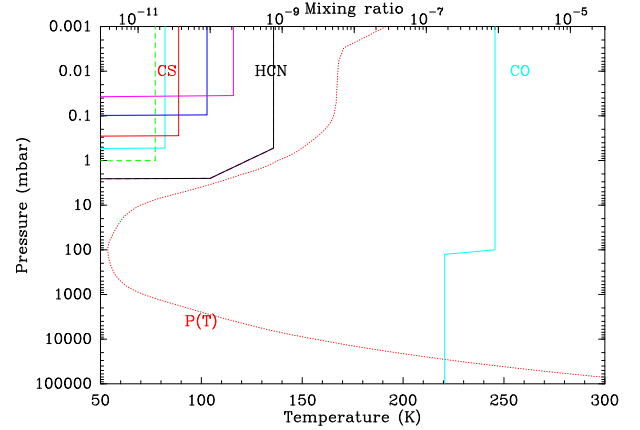


**Fig. 1.** ALMA individual observations of CS(7–6) at 342.8828503 GHz on Neptune at different dates over April 23–30, 2016, and averaged spectrum (*lower left panel*). Both horizontal and vertical polarisations are averaged. Observations are scaled in intensity to the conditions of April 30, 2016. The spectral resolution is  $\Delta\nu = 1.0$  MHz.



**Fig. 2.** Models of the CS(7–6) line, assuming a uniform mixing ratio distribution of CS at pressure levels lower than  $P_0$ , where  $P_0$  is varied from 0.03 to 1.0 mbar (cases A to E in Table 2). These models are compared with the averaged spectrum shown in Fig. 1.

and also used for Neptune (Lellouch et al. 2005, 2015). Opacity sources of CS, CO, and HCN were included (using line parameters from Pickett et al. 1998), also included was the collision-induced absorption opacity due to the main compounds of Neptune’s atmosphere,  $H_2$ – $H_2$ ,  $H_2$ –He, and  $H_2$ – $CH_4$ , using codes developed by Borysow et al. (1985, 1988), Borysow & Frommhold (1986). We adopted a He mole fraction of 0.149 (Burgdorf et al. 2003) and the  $CH_4$  vertical distribution of Lellouch et al. (2015). The Neptune thermal profile, taken from Lellouch et al. (2010), is shown in Fig. 3, along



**Fig. 3.** Neptune atmospheric model: thermal profile and vertical distributions of CO, HCN, and CS.

**Table 2.** Retrieved CS mixing ratio and column density assuming uniform profiles at pressure less than  $P_0$ .

Profile	$P_0$ (mbar)	Mixing ratio	Column ( $cm^{-2}$ )
A	0.03	$21.0^{+3.00}_{-3.00} \times 10^{-11}$	$2.0 \times 10^{12}$
B	0.1	$9.05^{+1.45}_{-1.35} \times 10^{-11}$	$2.2 \times 10^{12}$
C	0.3	$3.65^{+0.35}_{-0.30} \times 10^{-11}$	$2.5 \times 10^{12}$
D	0.5	$2.36^{+0.36}_{-0.36} \times 10^{-11}$	$3.1 \times 10^{12}$
E <sup>a</sup>	1.0	$1.73^{+0.25}_{-0.25} \times 10^{-11}$	$4.3 \times 10^{12}$

**Notes.** Mixing ratios include uncertainties in ALMA calibration and statistical uncertainties. (a) Modelled linewidth is too large.

with optimised CO and HCN profiles, retrieved from the high S/N observations of CO and HCN (not shown). As the CS pressure broadening coefficient ( $\gamma$ ) was not found in the literature, we nominally used the value of the HCN(4-3) line,  $\gamma = 0.121$   $cm^{-1}/atm$  with a temperature exponent  $n = 0.74$  (Landrain et al. 1997; Rohart et al. 2007).

The observed CS linewidth is consistent with broadening essentially due to the planet rotation, with at most minor Lorentz broadening. This precludes a measurement of the CS vertical profile, except for an upper limit on the pressure level at which it resides. Testing uniform CS profiles above a pressure level  $P_0$ , we inferred that the maximum allowable  $P_0$  is  $\sim 0.5$  mbar (Figs. 2 and 3). As the CS line is fully optically thin (e.g.  $\tau = 0.047$  for  $P_0 = 0.3$  mbar), the best determined quantity is the CS column density found to be in the range  $(2.0\text{--}3.1) \times 10^{12}$   $cm^{-2}$ , while the CS mixing ratio above  $P_0$  essentially scales as  $1/P_0$  and  $\sim 9 \times 10^{-11}$  for  $P_0 = 0.1$  mbar. Thus, CS is restricted to the upper stratosphere, similar to (and even higher than) HCN, and in sharp contrast with CO. We also tested a case with  $\gamma = 0.067$   $cm^{-1}/atm$  and  $n=0.64$  (i.e. the CO(3-2) broadening parameters, Sung 2004; Mantz et al. 2005). In this case the CS mixing ratio and column density is about 10% lower, and the maximum allowed  $P_0$  is  $\sim 0.9$  mbar. We thus conservatively conclude that the bulk of CS resides at pressures of less than 1 mbar.

## 4. Discussion

Our detection of CS in Neptune is the first unambiguous observation of a sulfur-bearing compound in a giant planet beyond Jupiter. Although  $H_2S$ , which has been established in Jupiter (Niemann et al. 1996), has been tentatively reported as

well in the deep atmospheres of Saturn, Uranus, and Neptune (Briggs & Sackett 1989; de Pater et al. 1991), this inference was based on the microwave spectrum in which the contributions of H<sub>2</sub>S and NH<sub>3</sub> are difficult to disentangle. At any rate, the putative presence of H<sub>2</sub>S in Neptune is almost certainly unrelated to that of CS. Indeed, any tropospheric H<sub>2</sub>S is expected to condense out near the 10 bar level, and therefore cannot serve as a source for CS at submillibar levels.

Prior to our measurements, and with the same motivation to confirm the comet impact scenario for CO, Iino et al. (2014) used the ASTE 10 m telescope to target a suite of sulfur-bearing species (CS, OCS, H<sub>2</sub>CS, H<sub>2</sub>S, SO, C<sub>3</sub>S, and SO<sub>2</sub>). None of these species was detected, but stringent upper limits were obtained. For CS, assuming uniform mixing at  $p < 6$  mbar, Iino et al. (2014) obtained a CS mixing ratio upper limit of  $1.8 \times 10^{-10}$ . Using for reference the CO mixing ratio measured by Hesman et al. (2007) for this pressure range (i.e.  $2.2 \times 10^{-6}$ ), this implied a CS/CO upper limit of  $8 \times 10^{-5}$ .

As detailed above, the CS stratospheric mixing ratio is  $2.4 \times 10^{-11}$  assuming that CS is present only at pressures lower than 0.5 mbar, and  $\sim 4$  times higher if CS is restricted to  $p < 0.1$  mbar. The most recent analyses of CO line profiles (combining data from *Herschel*/SPIRE and ground-based observations, Moreno et al., in prep.) indicate a CO tropospheric (resp. stratospheric) mixing ratio of  $0.2 \times 10^{-6}$  (resp.  $1.0 \times 10^{-6}$ ), indicative of a stratospheric excess of  $0.8 \times 10^{-6}$  of external origin that extends all the way down to 100 mbar. The above values for CS imply a CS/CO volume ratio of  $(3\text{--}11) \times 10^{-5}$  in the region where both species are present. However, as the external CO component extends much deeper than CS, a more relevant figure can be obtained by using the ratio of the column densities (or total masses) instead of the volume mixing ratios. The “external CO” column density ( $0.8 \times 10^{-6}$  mixing ratio above 100 mbar) is approximately  $1.8 \times 10^{19}$  cm<sup>-2</sup>. This gives a CS/CO column ratio of  $1.4 \times 10^{-7}$ .

As already noted by Iino et al. (2014), Neptune's stratospheric CS/CO ratio is much lower than Jupiter's. Based on data taken within 1 day after an impact, Lellouch et al. (1995) and Lellouch (1996) determined CS/CO =  $3 \times 10^{-3}$  by mass. The CS/CO ratio was found to be higher ( $\sim 0.01$  by mass)  $\sim 1$  week after the impacts (Moreno et al. 2001) and continued to increase to reach  $\sim 0.036$  by mass in the 1995–1998 time frame (Moreno et al. 2003). This evolution pointed to a likely photochemical build-up of CS from S<sub>2</sub> (Moses 1996), which was heavily produced from the impacts (Noll et al. 1995). The latter (i.e. long-term) value for CS/CO is probably more relevant for a comparison with Neptune. Thus, it appears that the CS/CO in Neptune's submillibar region is at least 200–1000 times smaller than in Jupiter's, and even much lower ( $\sim 1.6 \times 10^5$  times) if Neptune's CS/CO ratio is considered in terms of column density.

At face value this might be taken against the comet-impact source scenario for Neptune's CS (and CO and HCN). However, based on an unpublished monitoring of CO, CS, and HCN in Jupiter over 1996–2006, Moreno & Marten (2006) found that during that period, the total CO, CS, and HCN masses underwent global decreases by factors of  $\sim 6$ ,  $\sim 12$ , and  $\sim 8$ , respectively. For CS, a similar mass decrease of a factor of 11 was reported by Iino et al. (2016), but over 1996–2013. While the uncertainties on these figures are too large (typically a factor 2) to confidently conclude, they leave open the possibility that the evolution timescales for the different species are different. For example, adopting these nominal loss factors for CO, CS, and HCN over a ten-year interval leads to e-folding decay timescales of 5.4, 4.0, and 4.6 yr for CO, CS, and HCN. The

assumption of exponential decay is not proven, but reasonable for a unimolecular loss reaction such as sticking or a bimolecular one in which the other reactant is much more abundant. Although the direct applicability to Neptune is obviously hazardous to say the least, extrapolating this over a 200-yr timespan would lead to a depletion which is  $\sim 3 \times 10^5$  times stronger for CS than for CO. This effect is comparable to what is required to account for the much lower CS/CO mass ratio for Neptune (by a factor of  $\sim 1.6 \times 10^5$ ) compared to Jupiter. We also note that the same extrapolation over 200 yr would lead to a depletion of CS/HCN by 520, qualitatively matching the higher HCN/CS mass ratio on Neptune compared to Jupiter (by a factor of  $\sim 40$ , based on Fig. 3).

The other striking difference between Jupiter and Neptune is in the relative vertical distributions of CS and CO. In Jupiter, the two-decade monitoring indicates that the three SL9 related species, CO, CS, and HCN, have diffused downward from an initial deposition level of 0.1 mbar (Lellouch et al. 1995) to  $\sim 0.3$  mbar between 1998–2006 (Moreno et al. 2003; Moreno & Marten 2006) and appear to be roughly vertically co-located. In Neptune, CS is exclusively present in the submillibar level, while external CO invades the lower stratosphere, down to 6 mbar, 20 mbar, 2.5–100 mbar, and 100 mbar according respectively to Hesman et al. (2007), Lellouch et al. (2005), Luszcz-Cook & de Pater (2013, Fig. 19), and Moreno et al. (in prep.). Therefore, if CS was injected/produced following the same comet impact that is invoked for CO, it has not followed the same time evolution on the long term. The same is true for Neptune's HCN which –even if deposited by an ancient impact– is currently present at pressures lower than  $\sim 3$  mbar only, due to condensation. Carbon sulfide is stable (e.g. against photolysis) in the gas phase, but is a highly reactive species that polymerises on walls during laboratory experiments (Moltzen et al. 1988) and therefore has no measurable vapour pressure. Thus, while it does not “condense” in the traditional sense, CS can probably easily stick to aerosol surfaces or otherwise cluster together to form CS polymers; heterogeneous reactions on grain surfaces may also convert some of the CS to H<sub>2</sub>CS or to CH<sub>3</sub>CS via hydrogenation (J.-C. Loison, priv. comm.). Neptune's hazes extend up to the mbar level, fed by the condensation of species such as HCN, H<sub>2</sub>O, C<sub>4</sub>H<sub>2</sub>, and other heavier hydrocarbons. In the framework of the “old impact” hypothesis, this could be the reason for the vastly different vertical distributions of CO, HCN, and CS in Neptune. In Jupiter, the three species are still co-located given their recent deposition, but may vertically segregate in the longer term due to the same processes.

The comet impact hypothesis of Lellouch et al. (2005) seems to be a viable explanation for CO, HCN, and CS in Neptune. Lellouch et al. (2005) estimated a 200-yr impact age by equating it to the timescale ( $\tau = 2H^2/K$ ) required for diffusion down to the 20 mbar level, using an eddy K of 3000 cm<sup>2</sup> s<sup>-1</sup> at 20 mbar and an atmospheric scale height  $H = 30$  km, and inferred the diameter of the impactor to be  $D = 2$  km based on the amount of external CO they determined. In a slightly different formulation and considering lower values for K (e.g. 200–2000 cm<sup>2</sup> s<sup>-1</sup> at the tropopause), Luszcz-Cook & de Pater (2013) inferred timescales of 100–1000 yr. Furthermore, the most recent determinations of the CO profile (Luszcz-Cook & de Pater 2013; Moreno et al., in prep.) indicate a  $0.8 \times 10^{-6}$  CO stratospheric excess down to the  $\sim 100$  mbar level. As this represents 8 times more external CO than reported by Lellouch et al. (2005) (a  $0.5 \times 10^{-6}$  excess down to 20 mbar), we conclude that the single-impact scenario might have to be revised in favour of a slightly older – up to 1000 yr ago – but even bigger ( $D = 4$  km) comet impact.

Following the approach used by [Bézard et al. \(2002\)](#) to determine the origin of the non-SL9 CO stratospheric component on Jupiter, [Luszcz-Cook & de Pater \(2013\)](#) found that a permanent influx of small (sub-km size) comets is another possibility for Neptune's CO, perhaps preferable to the single, large-comet explanation given the statistical rarity of large impacts. We note, however, that the CO equivalent flux into Neptune is likely to be in excess of  $10^8 \text{ cm}^2 \text{ s}^{-1}$  ( $8 \times 10^8 \text{ cm}^2 \text{ s}^{-1}$  from scaling the [Lellouch et al. \(2005\)](#) flux by 8, and  $(0.5\text{--}20) \times 10^8 \text{ cm}^2 \text{ s}^{-1}$  according to the [Luszcz-Cook & de Pater \(2013\)](#) physical models), while the non-SL9 stratospheric CO on Jupiter requires a time-averaged flux of only  $(1.5\text{--}10) \times 10^6 \text{ cm}^2 \text{ s}^{-1}$  ([Bézard et al. 2002](#)). Cometary impacts have also been advocated as being responsible for the observed stratospheric CO in Saturn ([Cavalié et al. 2010](#)) and Uranus ([Cavalié et al. 2014](#)). While these two papers have considered the comet impact hypothesis in its “single large” impact version, a distribution of small comets is probably a viable solution for these two planets as well. But the very large external CO mixing ratio ( $0.8 \times 10^{-6}$ ) for Neptune, compared to  $(3\text{--}10) \times 10^{-9}$  for Jupiter ([Bézard et al. 2002](#)),  $(1\text{--}3) \times 10^{-8}$  for Saturn ([Cavalié et al. 2010](#)), and  $(0.7\text{--}1) \times 10^{-8}$  for Uranus ([Cavalié et al. 2014](#)), combined with the fact that statistical ecliptic comet impact rates ([Levison et al. 2000](#)) are not particularly enhanced on Neptune compared to other giant planets (e.g. no more than a factor of 2 more frequent than on Uranus), tends to argue for a unique  $\sim 4$  km comet impact on Neptune. Such an event may seem unlikely, but based on the [Zahnle et al. \(2003\)](#) cratering rates, we estimated that it occurs every 3400–4100 yr.

Alternate hypotheses do not seem easy to imagine. [Lellouch et al. \(1994\)](#) proposed that Triton's escaping atmosphere could be a source for the nitrogen present in Neptune's HCN and found this mechanism to require a total nitrogen infall rate of  $5 \times 10^{24} \text{ N atoms s}^{-1}$ , in plausible agreement with estimates of Triton's escape rate, and of ionisation and transport rates of nitrogen to Neptune. However, the process is likely to be irrelevant for CS, given the lack of any detected atmospheric or surface sulfur-bearing species on Triton (see e.g. [Grundy et al. 2010](#), for a high S/N near-IR spectrum of Triton).

Given the elemental abundance composition of cometary dust ([Jessberger & Kissel 1991](#), i.e. O/S = 12.4), interplanetary dust particles also represent a source of oxygen and other volatiles, with an O/S ratio broadly consistent with solar composition (chondritic O/S = 8.5, photospheric solar O/S = 31.7; [Lodders 2003](#)). Thus, in terms of chemistry, interplanetary dust particles (IDP) may be equally viable sources of sulfur as cometary impacts. However, the oxygen influx required to explain the abundances of CO on Neptune is vastly different ( $\geq 10^8 \text{ cm}^2 \text{ s}^{-1}$ , [Lellouch et al. 2005](#); [Luszcz-Cook & de Pater 2013](#); [Hesman et al. 2007](#)) from that needed to account for the CO, H<sub>2</sub>O, and CO<sub>2</sub> ( $\sim (2\text{--}3) \times 10^5 \text{ cm}^2 \text{ s}^{-1}$  [Cavalié et al. 2014](#); [Orton et al. 2014](#)) on Uranus; instead, based on a dynamical interplanetary dust model, [Poppe \(2016\)](#) calculated an IDP oxygen flux only  $\sim 5$  times larger than into the atmosphere of Uranus. In fact, [Poppe \(2016\)](#) and [Moses & Poppe \(2017\)](#) find that a flux of IDPs is sufficient to account for the abundances of all three oxygen-bearing species on Uranus, as well as H<sub>2</sub>O on Neptune (and CO<sub>2</sub>, if CO is prescribed), but fails to explain Neptune's CO abundance. Based on their calculation of dust ablation, [Moses & Poppe \(2017\)](#) conclude that the bulk of the stratospheric oxygen on Neptune originates from a cometary impact and not from IDPs.

In summary, the CS detection, along with recent analyses of the CO profile, reinforces the presumption of a large comet

impact into Neptune up to  $\sim 1000$  yr ago that delivered CO, CS, and HCN at the same time. The three species may have been produced in proportions that were initially similar to the SL9 events, but relative amounts were progressively altered by the preferential condensation/sticking of HCN near 3 mbar, and CS polymerisation or other chemical loss at submillibar levels.

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## References

- Bézard, B., Lellouch, E., Strobel, D., Maillard, J.-P., & Drossart, P. 2002, *Icarus*, **159**, 95
- Borysow, A., & Frommhold, L. 1986, *ApJ*, **304**, 849
- Borysow, J., Trafton, L., Frommhold, L., & Birnbaum, G. 1985, *ApJ*, **296**, 644
- Borysow, J., Frommhold, L., & Birnbaum, G. 1988, *ApJ*, **326**, 509
- Briggs, F. H., & Sackett, P. D. 1989, *Icarus*, **80**, 77
- Burgdorf, M., Orton, G. S., Davis, G. R., et al. 2003, *Icarus*, **164**, 244
- Cavalié, T., Billebaud, F., Dobrijevic, M., et al. 2009, *Icarus*, **203**, 531
- Cavalié, T., Hartogh, P., Billebaud, F., et al. 2010, *A&A*, **510**, A88
- Cavalié, T., Moreno, R., Lellouch, E., et al. 2014, *A&A*, **562**, A33
- de Pater, I., Romani, P. N., & Atreya, S. K. 1991, *Icarus*, **91**, 220
- Feuchtgruber, H., Lellouch, E., de Graauw, T., et al. 1997, *Nature*, **389**, 159
- Grundy, W. M., Young, L. A., Stansberry, J. A., et al. 2010, *Icarus*, **205**, 594
- Hesman, B. E., Davis, G. R., Matthews, H. E., & Orton, G. S. 2007, *Icarus*, **186**, 342
- Iino, T., Mizuno, A., Nakajima, T., et al. 2014, *Planet. Space Sci.*, **104**, 211
- Iino, T., Ohshima, H., Hirahara, Y., Takahashi, T., & Tsukagoshi, T. 2016, *AJ*, **152**, 179
- Jessberger, E. K., & Kissel, J. 1991, in *Comets in the post-Halley era*, eds. R. L. Newburn, Jr., M. Neugebauer, & J. Rahe, *Astrophys. Space Sci. Lib.*, **167**, IAU Colloq. **116**, 1075
- Landrain, V., Blanquet, G., Lepère, M., Walrand, J., & Bouanich, J.-P. 1997, *J. Mol. Spectr.*, **182**, 184
- Lellouch, E. 1996, in *The Collision of Comet Shoemaker-Levy 9 and Jupiter*, eds. K. S. Noll, H. A. Weaver, & P. D. Feldman, *IAU Colloq.*, **156**, 213
- Lellouch, E., Romani, P. N., & Rosenqvist, J. 1994, *Icarus*, **108**, 112
- Lellouch, E., Paubert, G., Moreno, R., et al. 1995, *Nature*, **373**, 592
- Lellouch, E., Bézard, B., Moses, J. I., et al. 2002, *Icarus*, **159**, 112
- Lellouch, E., Moreno, R., & Paubert, G. 2005, *A&A*, **430**, L37
- Lellouch, E., Hartogh, P., Feuchtgruber, H., et al. 2010, *A&A*, **518**, L152
- Lellouch, E., Moreno, R., Orton, G. S., et al. 2015, *A&A*, **579**, A121
- Levison, H. F., Duncan, M. J., Zahnle, K., Holman, M., & Dones, L. 2000, *Icarus*, **143**, 415
- Lodders, K. 2003, *ApJ*, **591**, 1220
- Luszcz-Cook, S. H., & de Pater, I. 2013, *Icarus*, **222**, 379
- Mantz, A. W., Malathy Devi, V., Chris Benner, D., et al. 2005, *J. Mol. Struct.*, **742**, 99
- Marten, A., Gautier, D., Owen, T., et al. 1993, *ApJ*, **406**, 285
- Moltzen, E. K., Klabunde, K. J., & Senning, A. 1988, *Chem. Rev.*, **88**, 391
- Moreno, R., & Marten, A. 2006, *BAAS*, **38**, 497
- Moreno, R., Marten, A., Biraud, Y., et al. 2001, *Planet. Space Sci.*, **49**, 473
- Moreno, R., Marten, A., Matthews, H. E., & Biraud, Y. 2003, *Planet. Space Sci.*, **51**, 591
- Moses, J. I. 1996, in *The Collision of Comet Shoemaker-Levy 9 and Jupiter*, eds. K. S. Noll, H. A. Weaver, & P. D. Feldman, *IAU Colloq.*, **156**, 243
- Moses, J. I., & Poppe, A. R. 2017, *Icarus*, **297**, 33
- Niemann, H. B., Atreya, S. K., Carignan, G. R., et al. 1996, *Science*, **272**, 846
- Noll, K. S., McGrath, M. A., Trafton, L. M., et al. 1995, *Science*, **267**, 1307
- Orton, G. S., Moses, J. I., Fletcher, L. N., et al. 2014, *Icarus*, **243**, 471
- Pickett, H. M., Poynter, R. L., Cohen, E. A., et al. 1998, *J. Quant. Spectr. Rad. Transf.*, **60**, 883
- Poppe, A. R. 2016, *Icarus*, **264**, 369
- Rohart, F., Nguyen, L., Buldyreva, J., Colmont, J.-M., & Włodarczyk, G. 2007, *J. Mol. Spectr.*, **246**, 213
- Sung, K. 2004, *J. Quant. Spectr. Rad. Transf.*, **83**, 445
- Zahnle, K., Schenk, P., Levison, H., & Dones, L. 2003, *Icarus*, **163**, 263